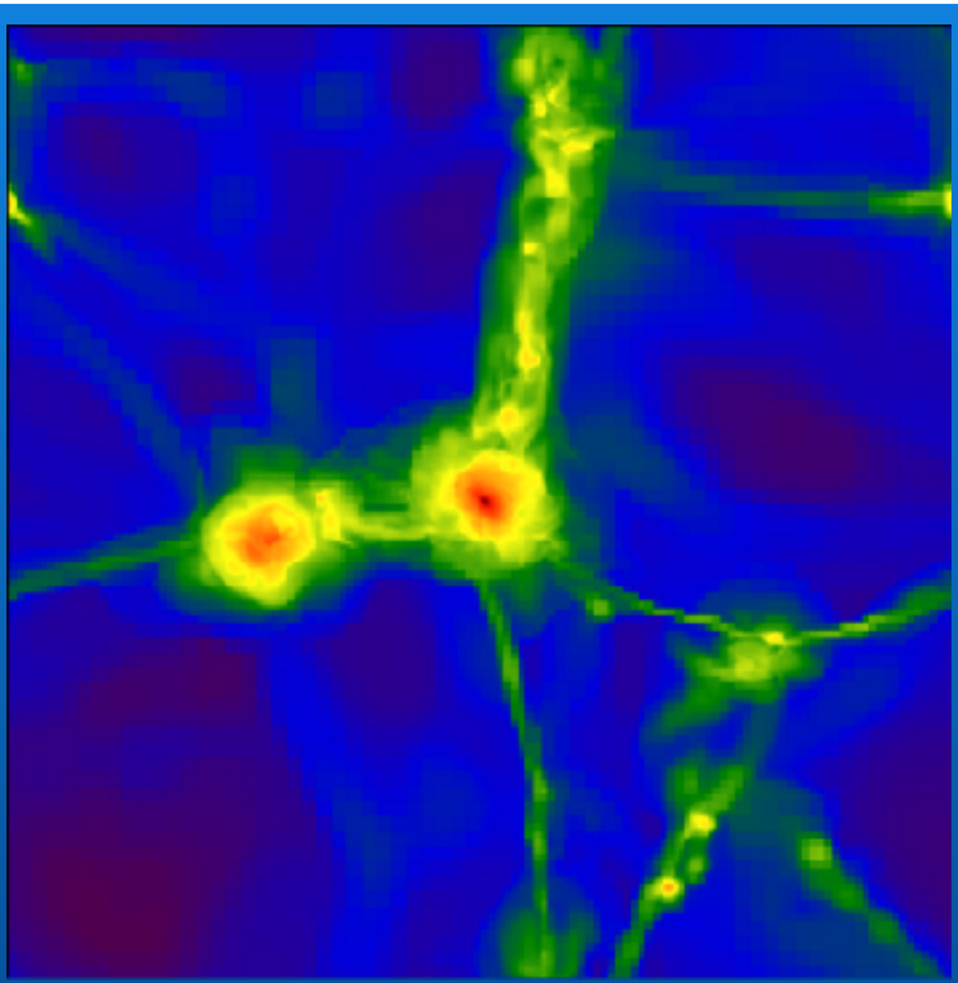
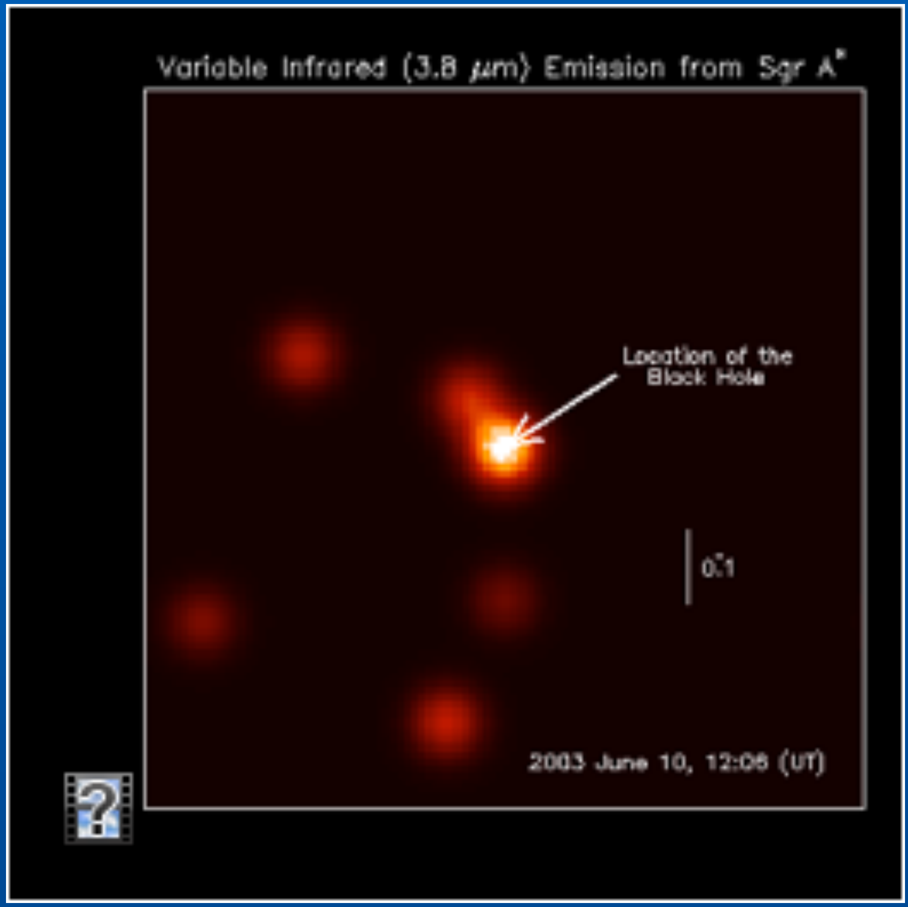
Unveiling the first black holes



Ricarte+; Pacucci+; Cappelluti+; Agarwal+; Haiman+; Tanaka+; Johnson+; Park+; Ricotti+; Khochfar+; Di Matteo+; Yoshida+ Schneider+Dubois+; Bournoud+; Ferrara+; Bromm+; Milosavljevic+; Ricotti+; Norman+; Omukai+; Inayoshi+; Regan+; Sijacki+; Loeb+; Habuzit+; Smith+; Volonteri+; Abel+; Wise+; Latif+; Whalen+; Trakhtenbrot+; Weigel+: Schawinski+;

Priyamvada Natarajan
Yale University

MULTI-WAVELENGTH DATA FOR ACTIVE & QUIESCENT BHs

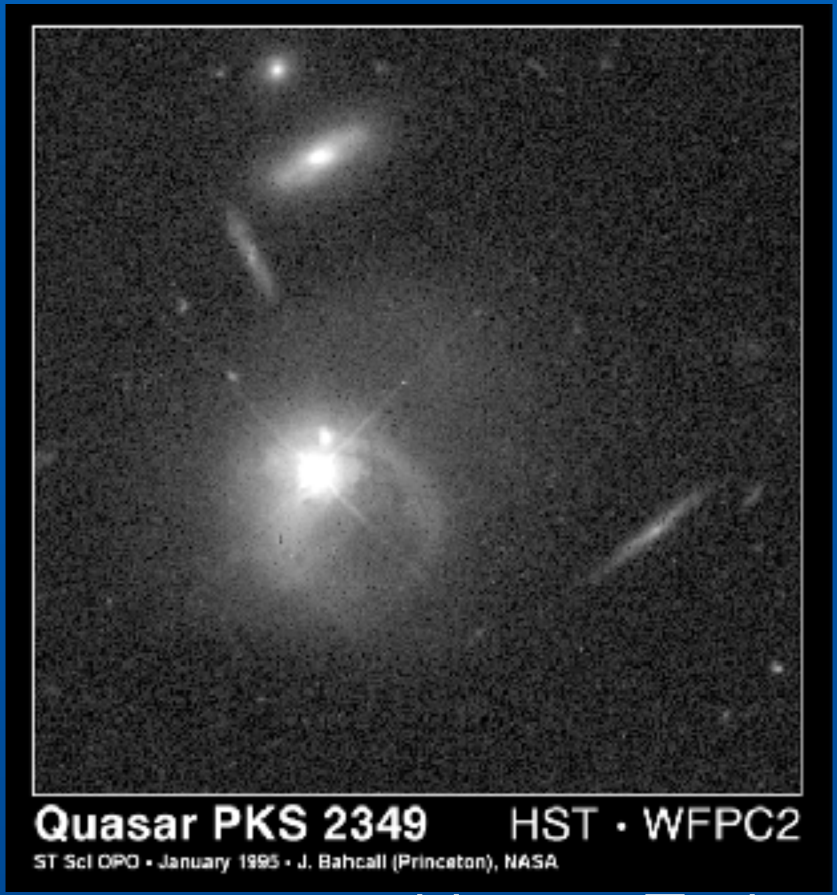


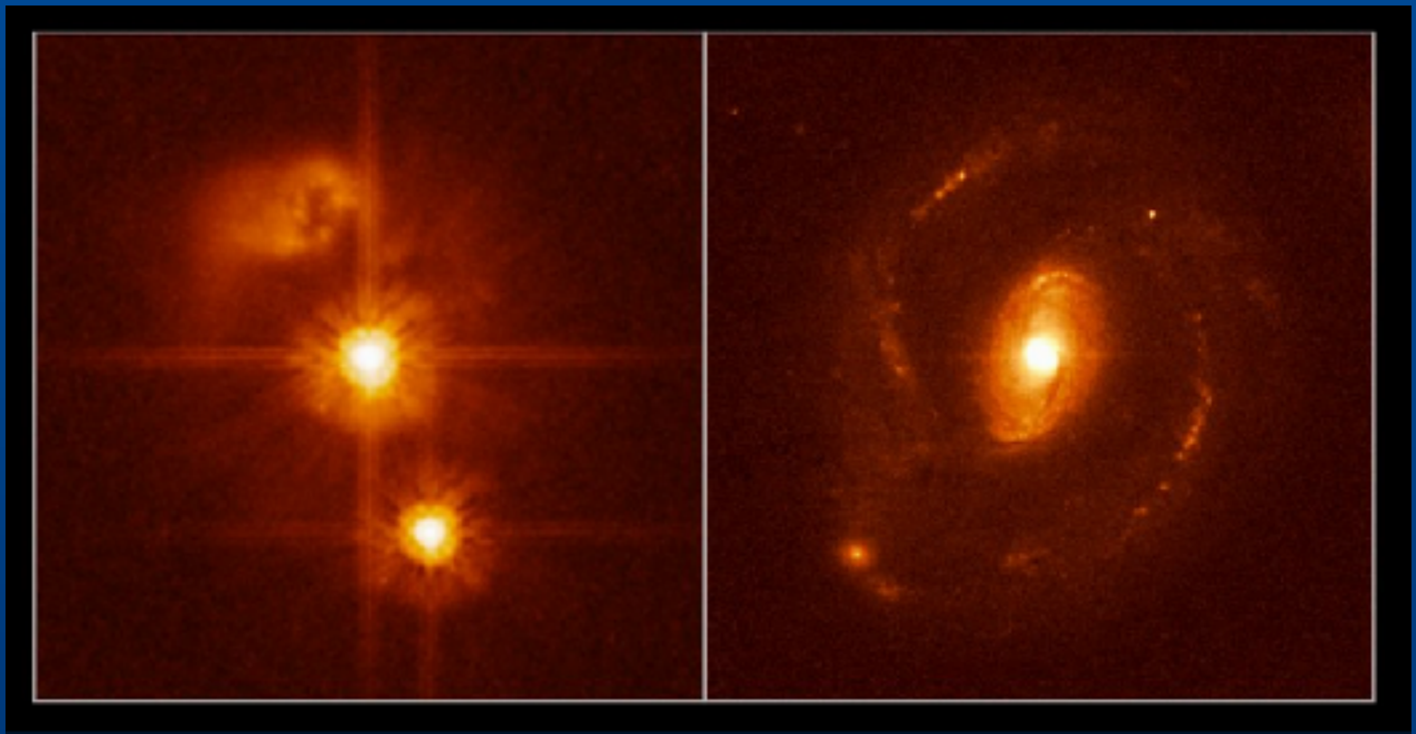
$M_{BH} \sim 10^{6\text{-}9}\ M_{sun}$

even $10^{10}\ M_{sun}$

z ~ 0 – 7

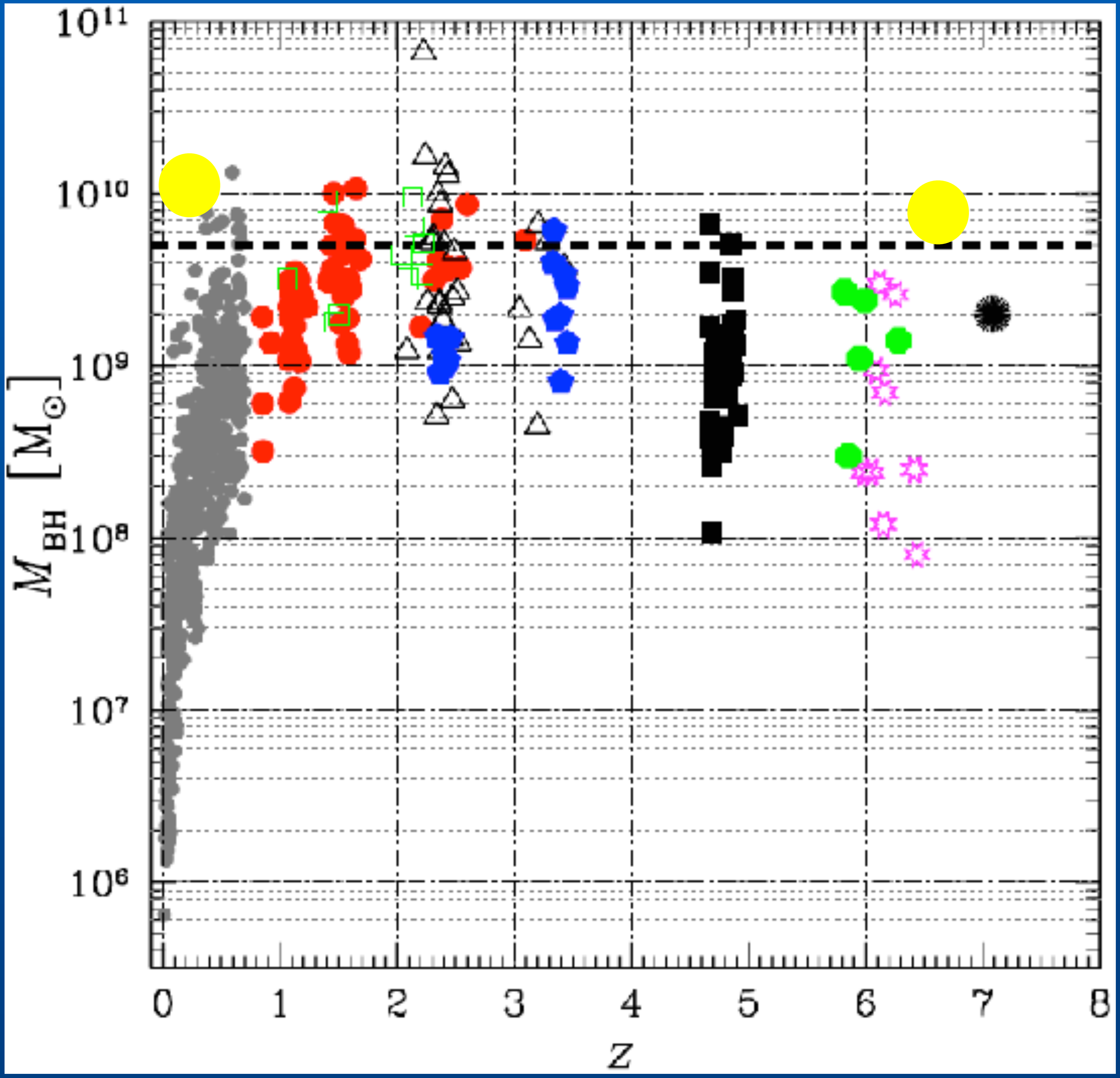
z=7
756 Myr after
the Big Bang

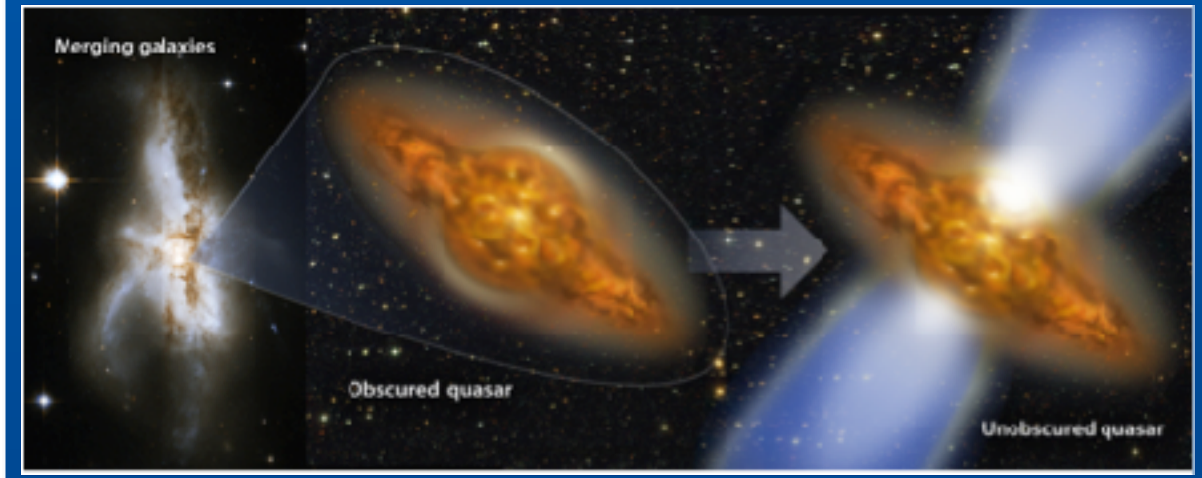




Urry+; Treister+; Scoville+; Sanders; Faber+;Wu+; Fan+; Ferguson+; Harrison+; Hasinger+; Comastri+; Gilli+

The Most Massive Black Holes in the Universe





Lauer+ 05, 06; Bernardi+ 06; PN & Treister 09; McConnell+ 11,12; PN & Volonteri 13
Marziani & Sulentic 12; Mortlock+ 14; Wu+ 2015; Kulier+15; Thomas+ 16

BHMF FOR BLQSOs FROM SDSS 1 < z < 4.5

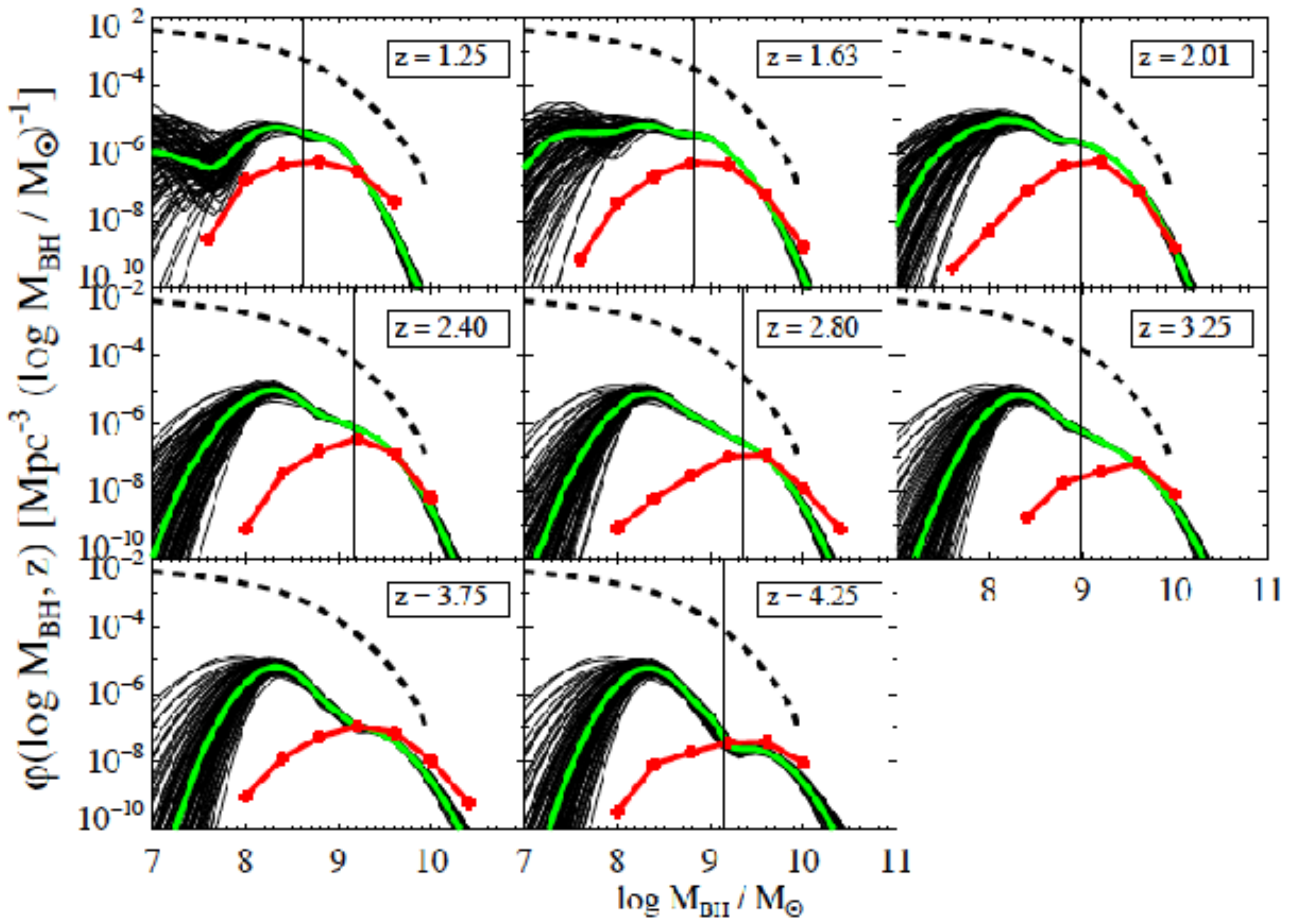


FIG. 3.— BLQSO BHMF (thin solid lines) obtained using our Bayesian approach, compared with the local BHMF fo all SMBHs (dashed line), and the BHMF from Vestergaard et al. (2008, solid red line with points); as in Figure 1, each thin solid line denotes a random draw of the BHMF from its probability distribution. The thick green line is the median of the BHMF random draws, and may be considered our 'best-fit' estimate. The vertical line marks the mass at which the SDSS DR3 sample becomes 10% complete.

Kelly, Shen+ 10; 11

HIGH-z QUASARS & THE TIMING CRUNCH TO ASSEMBLE SMBHs

Bright quasars host 10^9 - 10^{10} M_{sun} BHs

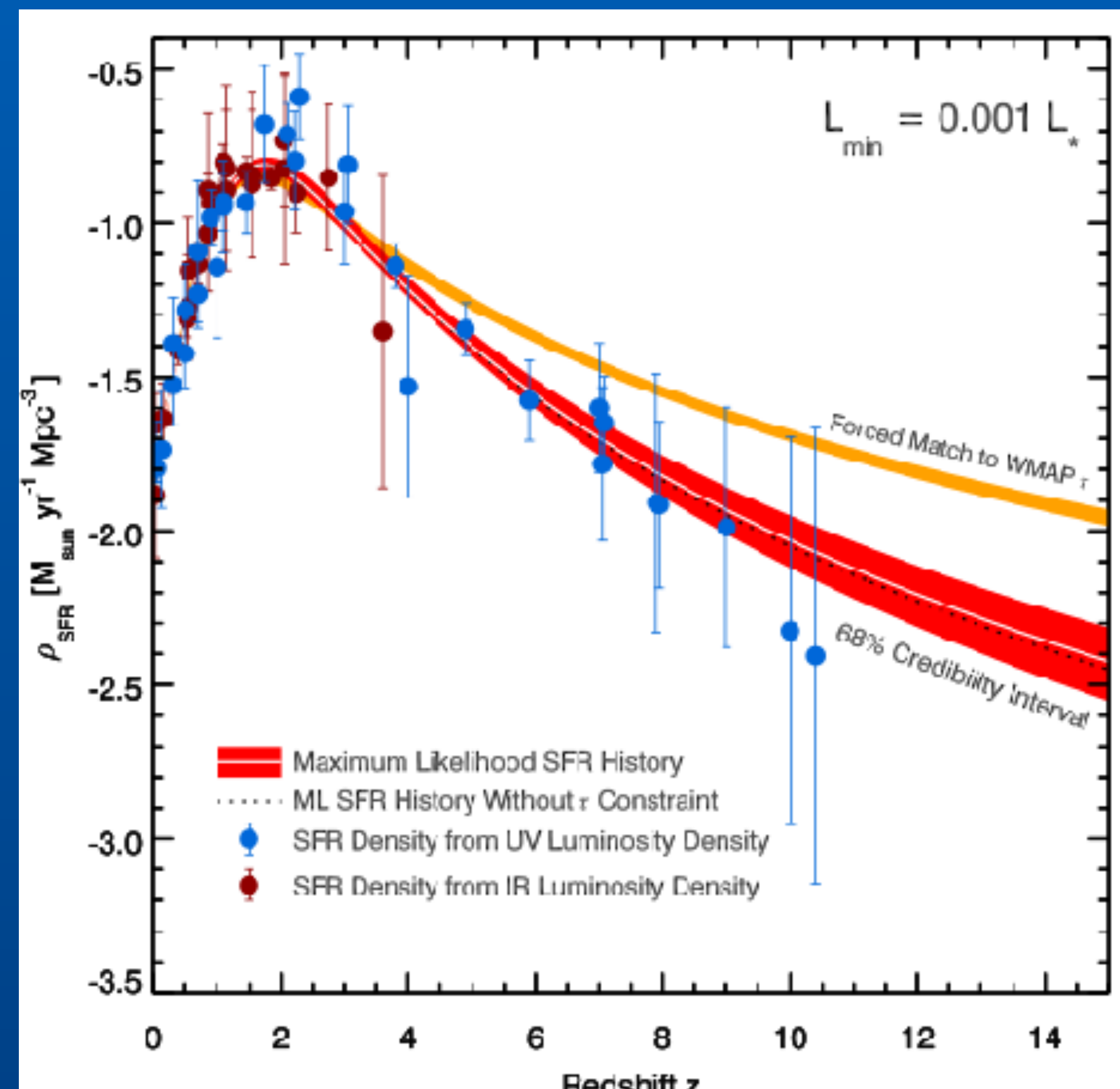
Age of the universe 1 Gyr

Eddington limit growth rate of mass

$$\frac{dM}{dt} = \frac{L_{acc}}{\eta c^2} < \frac{4\pi G M m_p}{\eta c \sigma_T}$$

$$M \le M_0 e^{\frac{t}{\tau}}$$

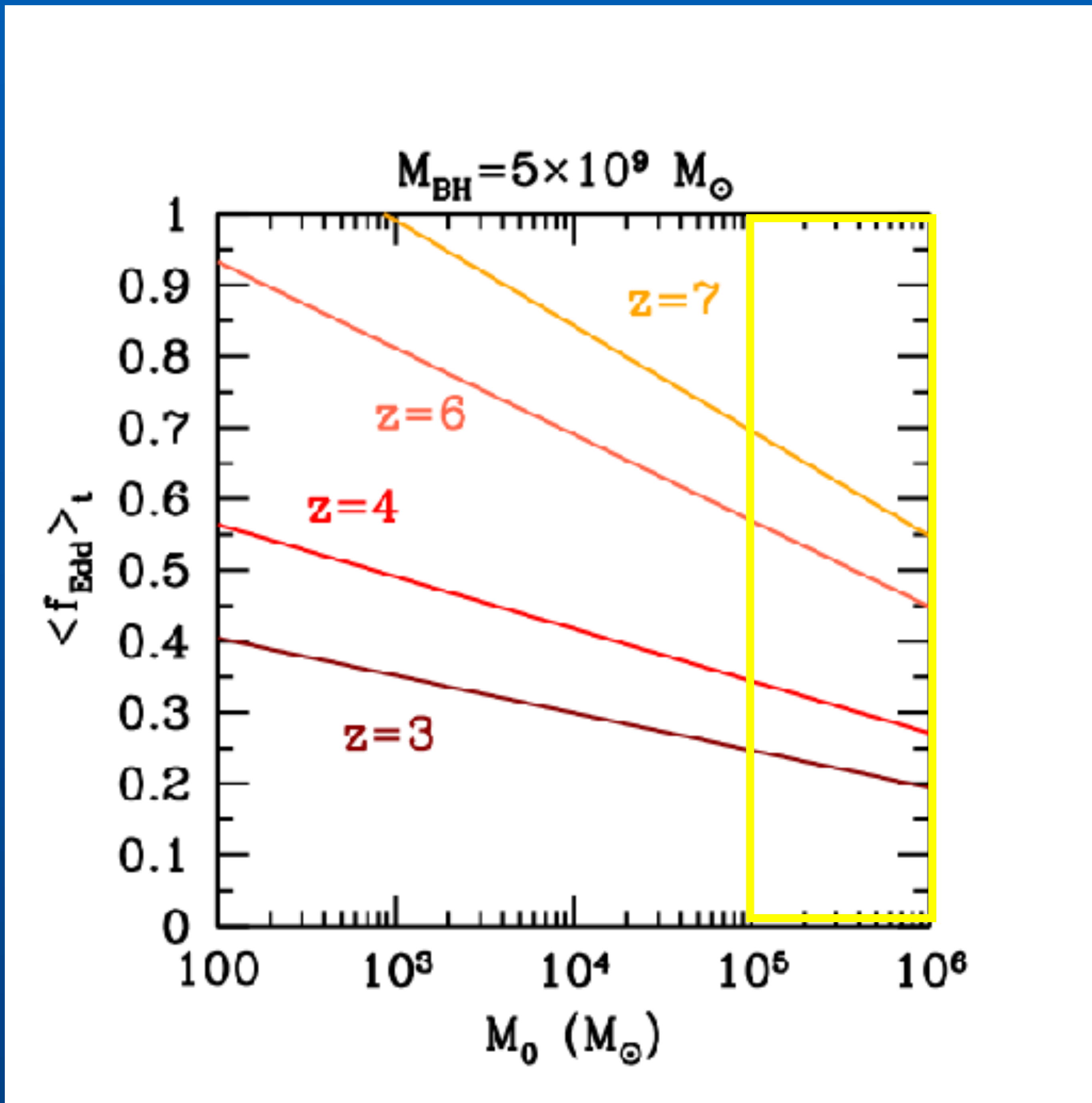
$$\tau = \frac{\eta c \sigma_T}{4\pi G m_p} \approx 5\times 10^7\, yr$$



LATEST PLANCK RESULTS: first stars form even later!

Wu+ 15; Robertson+ 15; Planck+ XIII 15

MASS GROWTH OF BH SEEDS: TIME CRUNCH



$$\langle f_{\rm Edd}\rangle_t = \frac{t_{\rm Edd}}{t_{\rm Hubble}(z)}\frac{\epsilon}{1-\epsilon}\ln\left(\frac{M_{\rm BH}}{M_0}\right).$$

$$L = \epsilon\,\dot{M}_{in}c^2 = f_{\rm Edd}L_{\rm Edd}c^2,$$

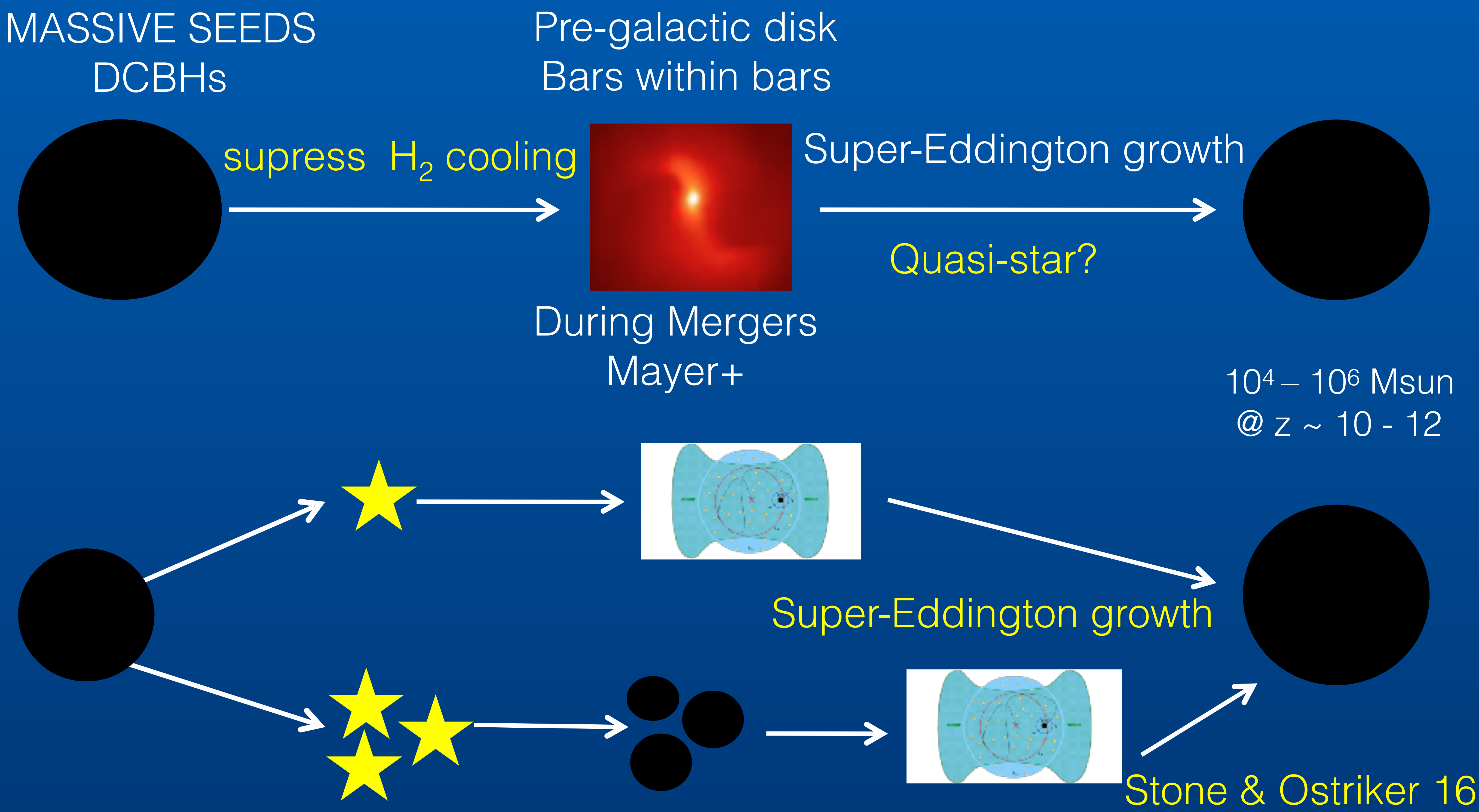
AGE OF THE UNIVERSE AT z = 7 [771 Myr]; z= 4[1.57 Gyr]; z= 3 [2.9 Gyr]

PN+ 17; Treister+ 13

CONDITIONS REQUIRED TO ASSEMBLE THE MOST MASSIVE BLACK HOLES AT ALL EPOCHS

- Massive initial BH seeds - rapid early growth - at high z (trade-off between obscured growth vs. massive initial seeds)
- Brief periods of super-Eddington accretion early on
- Periods of obscured growth during merger triggered phases

SYNOPSIS OF CURRENT VIEW ON BH SEEDS TO MAKE THE MOST MASSIVE HIGH-x BHs



LIGHT SEEDS

PopIII

~10^{1-2} M_{sun}

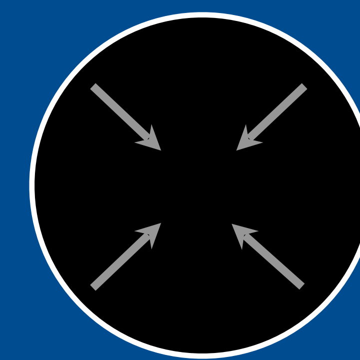
MASSIVE SEEDS

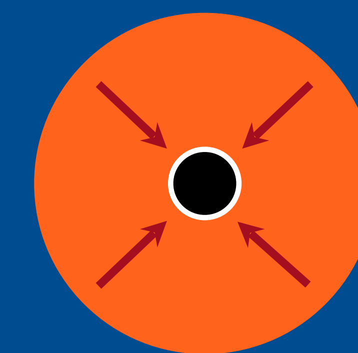
Direct collapse

Nuclear star cluster

~10^{3} M_{sun}

Supermassive Star





Quasi star

~10^{4-6} M_{sun}

Uncertainity in the masses of the first stars

A challenge to grow monster BHs seen by t < 2 Gyrs

New Planck results push first stars to later even ~550 Myrs after the Big Bang

In protogalaxies: need to avoid fragmentation and star formation, need to centrally concentrate mass

- Metal-free gas
- Prevent molecular H-cooling

First black holes in pre-galactic halos z = 20-30

$M_{BH} \sim 1 - 100\ M_{sun}$ **LIGHT SEEDS**

Pop III remnants : Simulations suggest that the first stars have a range of masses (Bromm+ 02 ; Abel+ 02; Abel+ 00; Alvarez+ 08; Hirano+ 14) Metal free Pop III stars leave remnant BHs

Supra-exponential early growth boost: Super-Eddington growth in nuclear star clusters at high-z (Alexander & PN 14)

$M_{BH} \sim 10^3 - 10^6\ M_{sun}$ **MASSIVE SEEDS**

Direct Collapse – efficient viscous transport, H2 cooling suppressed, Lyman-Werner radiation, formation of central concentration (Eisenstein & Loeb 95; Koushiappas+ 04)+ proper dynamical treatment of disk stability (Lodato & PN 06, 07)

Supermassive star (Haehnelt & Rees 93)

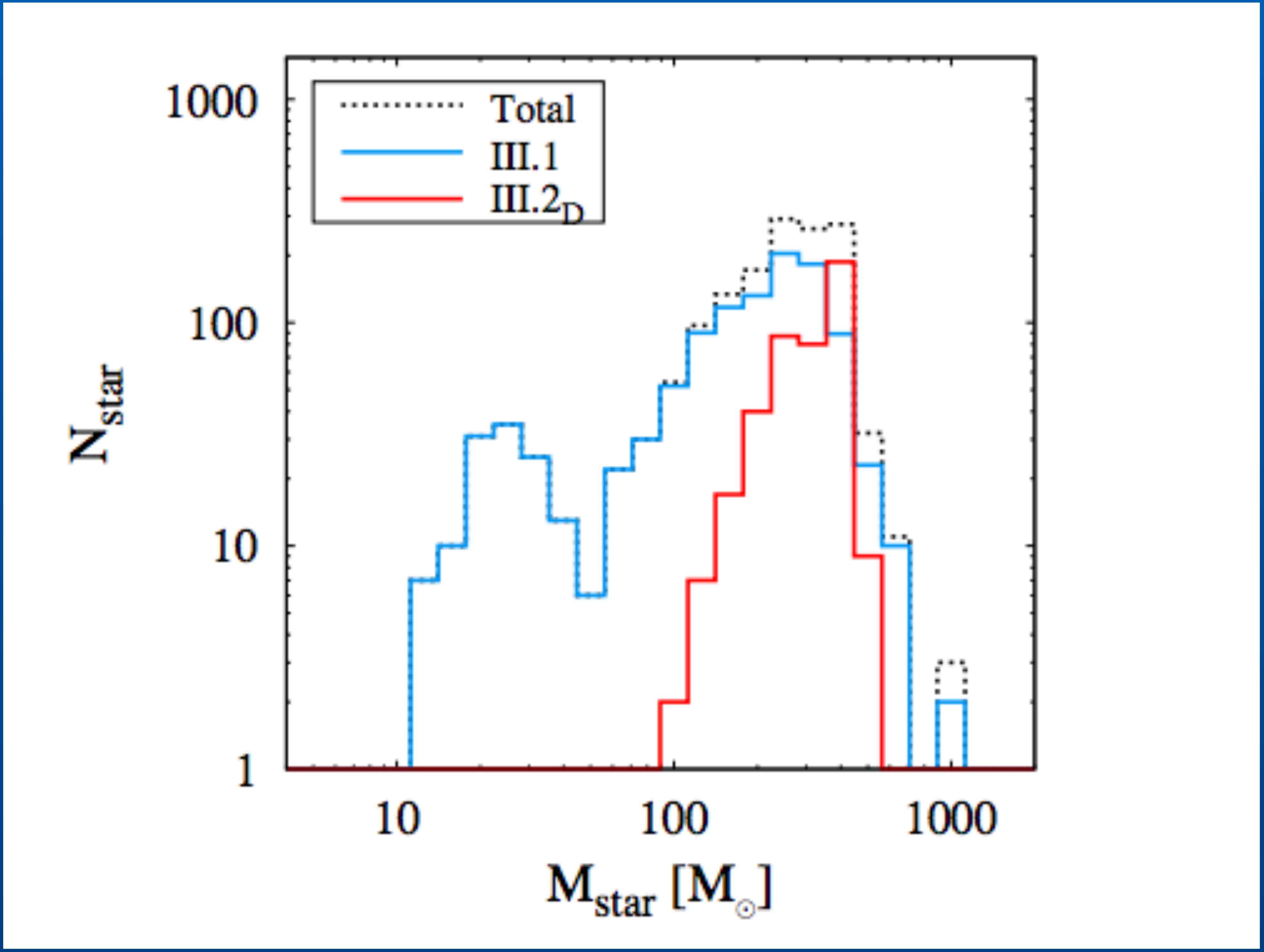
Quasi-star - Bar unstable self-gravitating gas + large quasi-star (Begelman 08; 10; 12)

PRIMORDIAL BLACK HOLES

Post-Inflation formation of black holes during the phase transition that ends expansion (Khlopov+ 07; 09)

Unresolved CIRB, XRB Excess? CMBR distortions, no stringent constraints

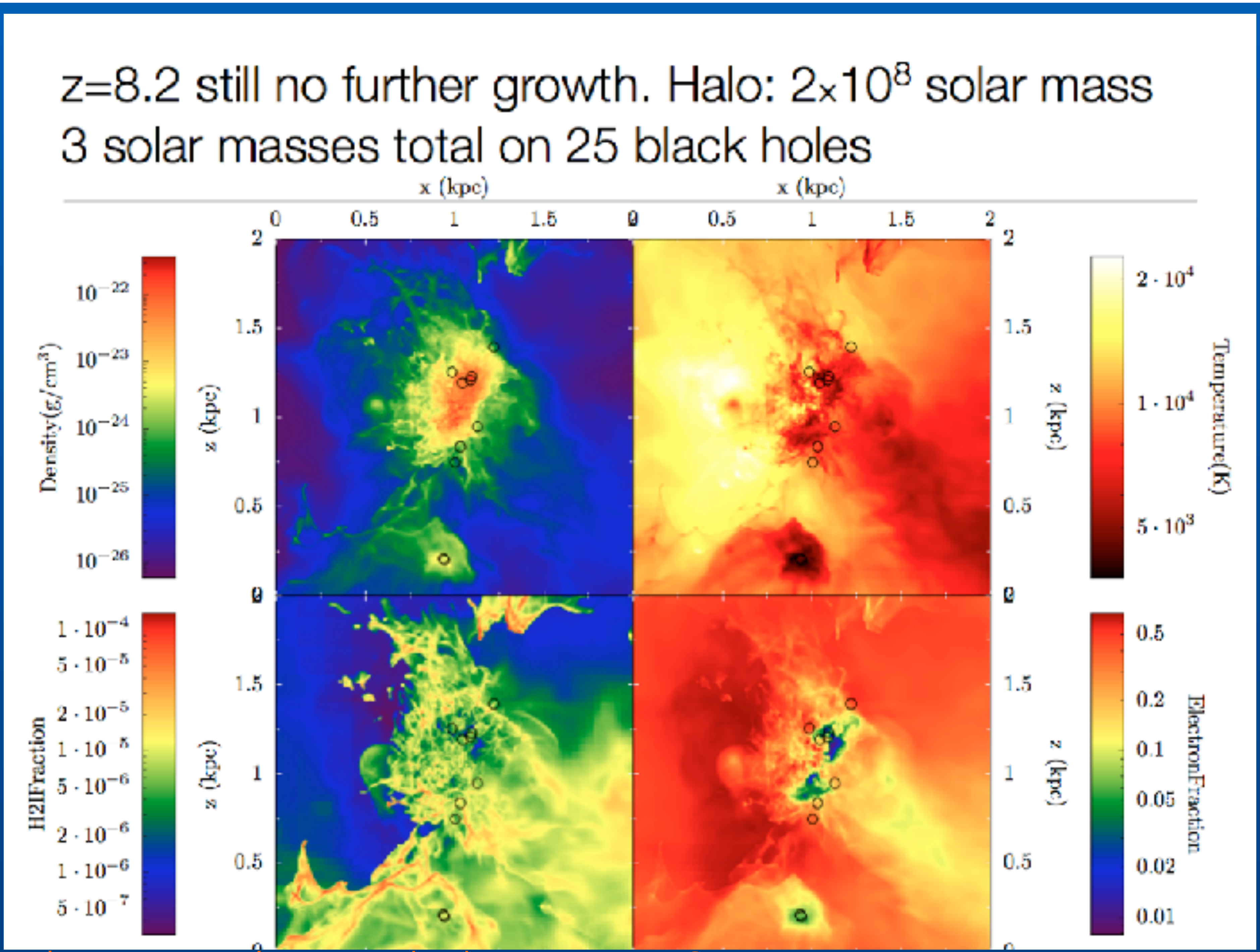
HOW MASSIVE ARE POP III STARS?



~ Mass distribution of Pop III stars formed at z = 30 -> 10

Hirano+ 15

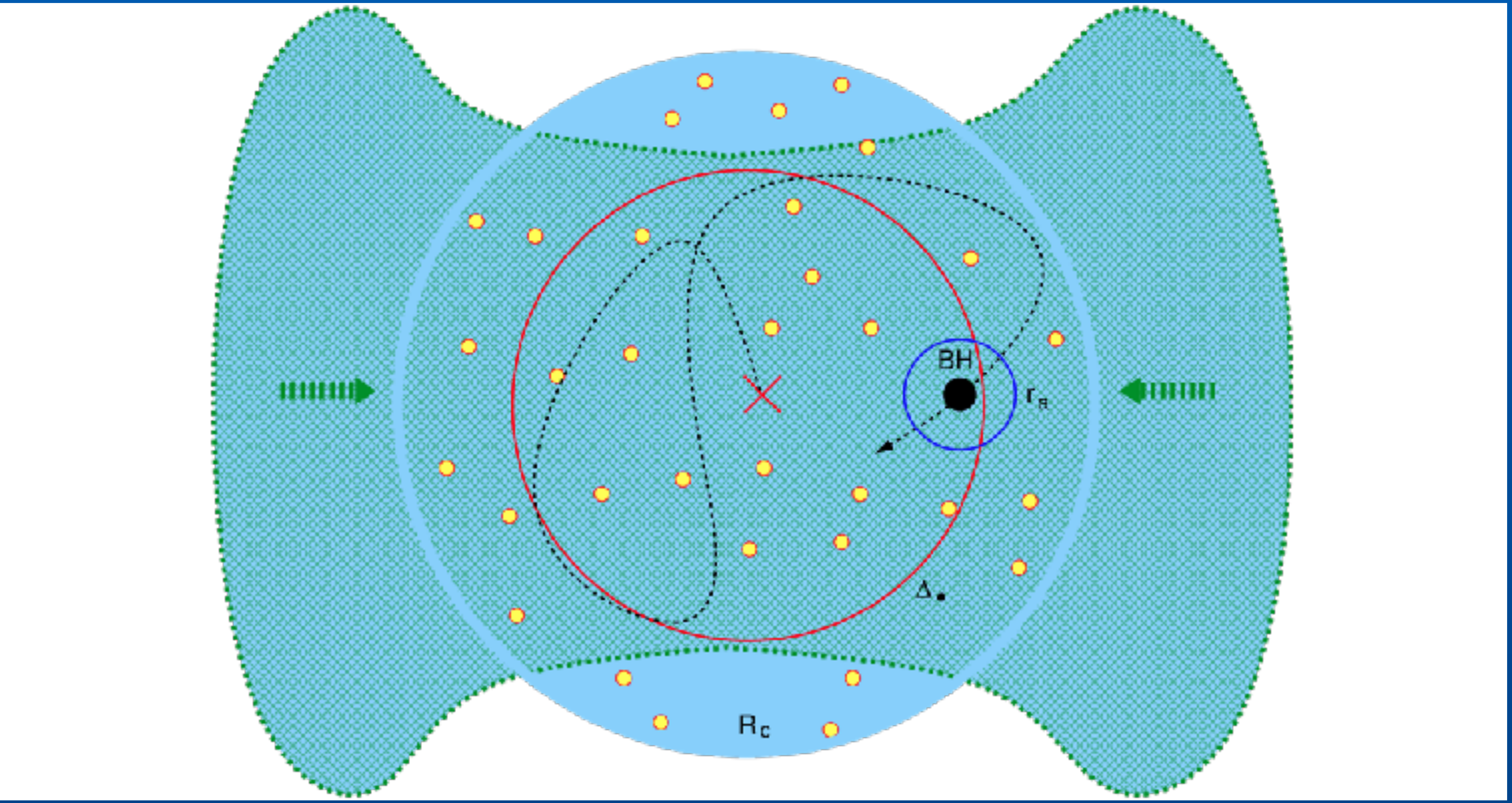
DO WE NEED MASSIVE BH SEEDS?
Tracking the fate of PopIII seeds in 2.5-sigma peaks



BHs simply not growing much down to z = 8
even when PopIII formation has ceased
BHs spend almost all their time in the wrong place
in 10^8 M_{sun} DM halos

Abel, Wise, Turk, Alvarez+; Stacy+

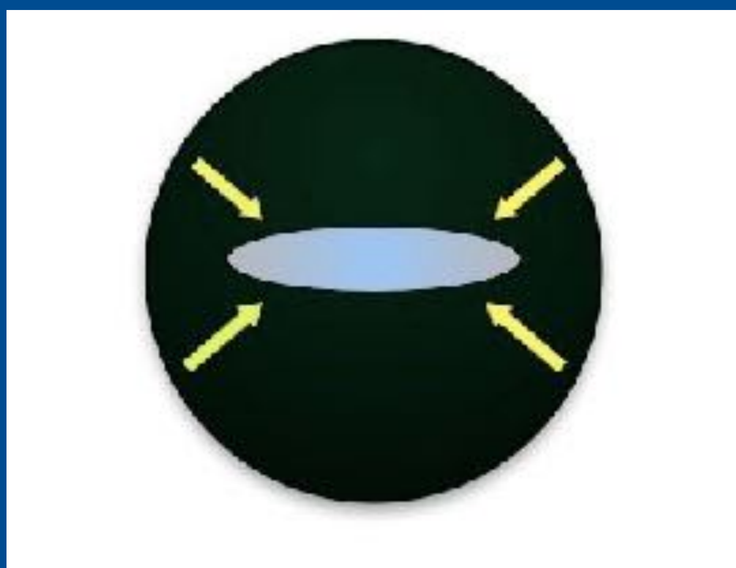
SUPER BOOSTING EARLY BH GROWTH



Circumventing the Eddington limit

Alexander & Natarajan 14

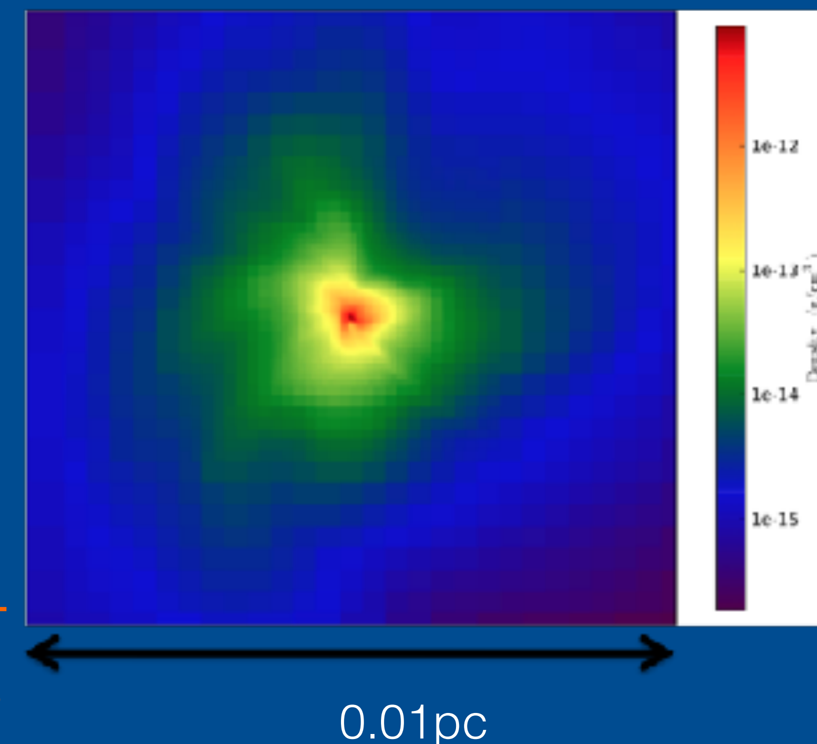
BH seed formation at high z

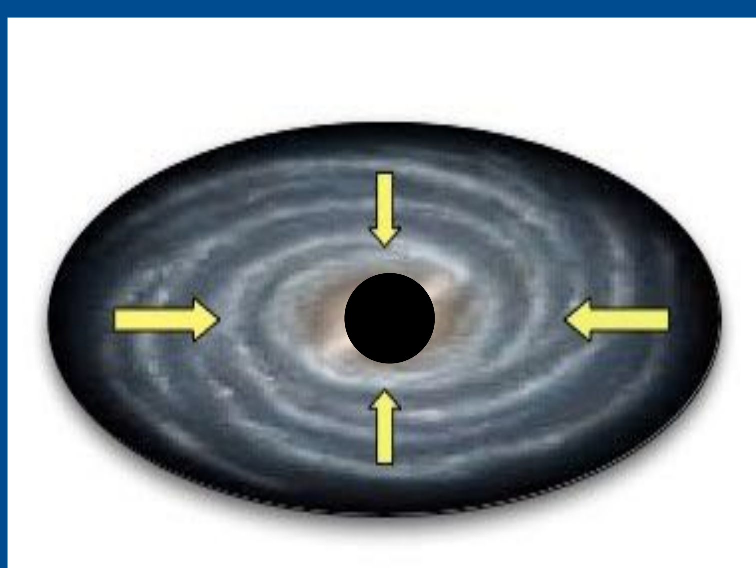


Baryons inside DM halo collapse and form a rotating pre-galactic disc

Disc becomes gravitationally unstable and accretes to the center

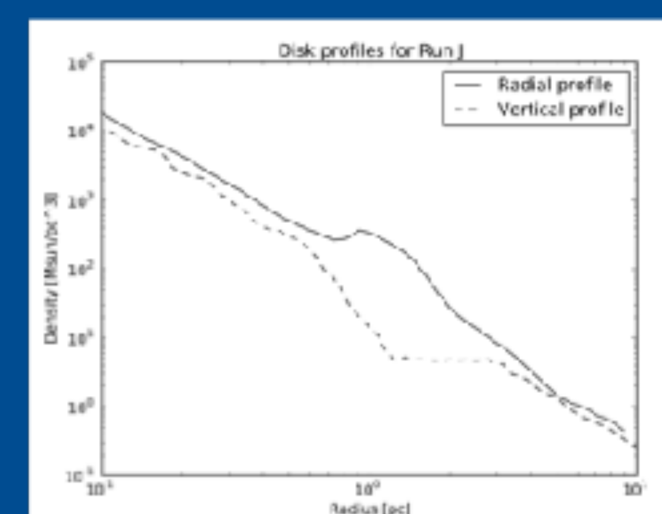
Angular mom of DM halo + Gas reservoir + dynamics (disc stability) + cooling





LYMAN WERNER PHOTONS FROM NEARBY SF HALO NEEDED TO DISSOCIATE MOL HYDROGEN TO PREVENT COOLING & FRAGMENTATION

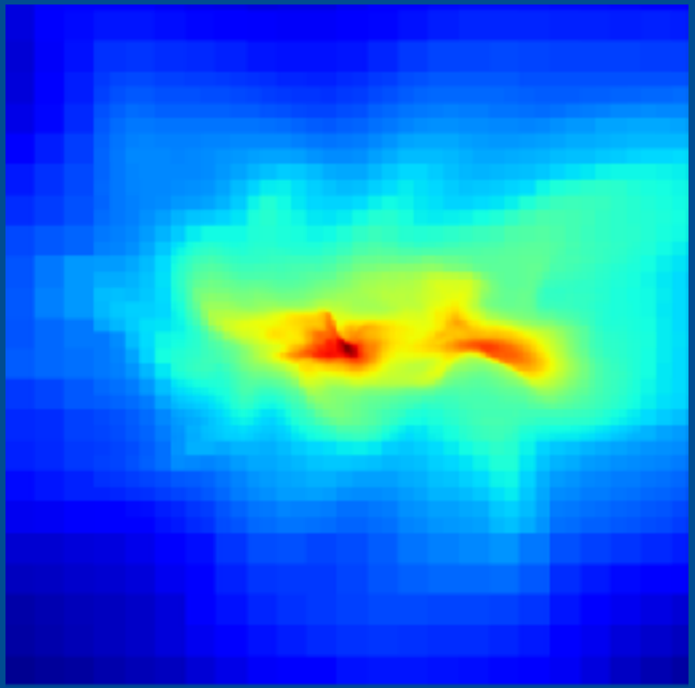
DCBHS FORM IN SATELLITE HALOS

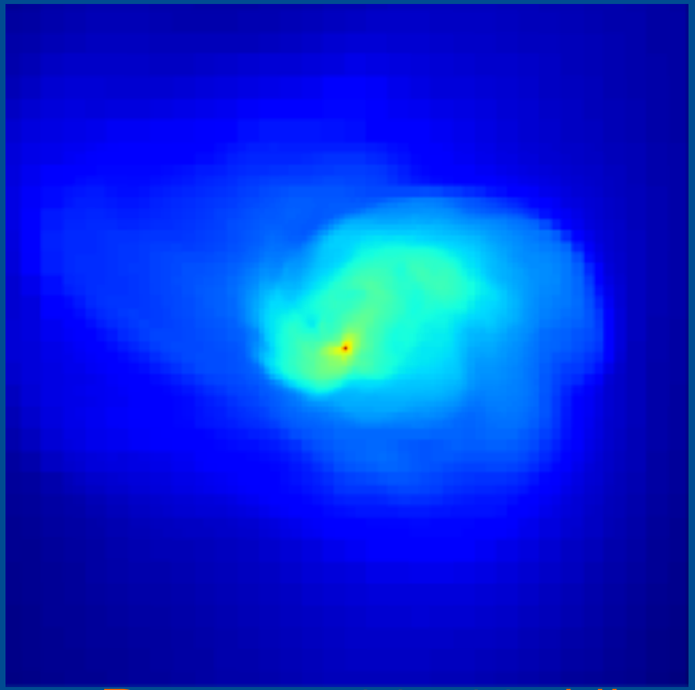


FINAL DCBH MASS

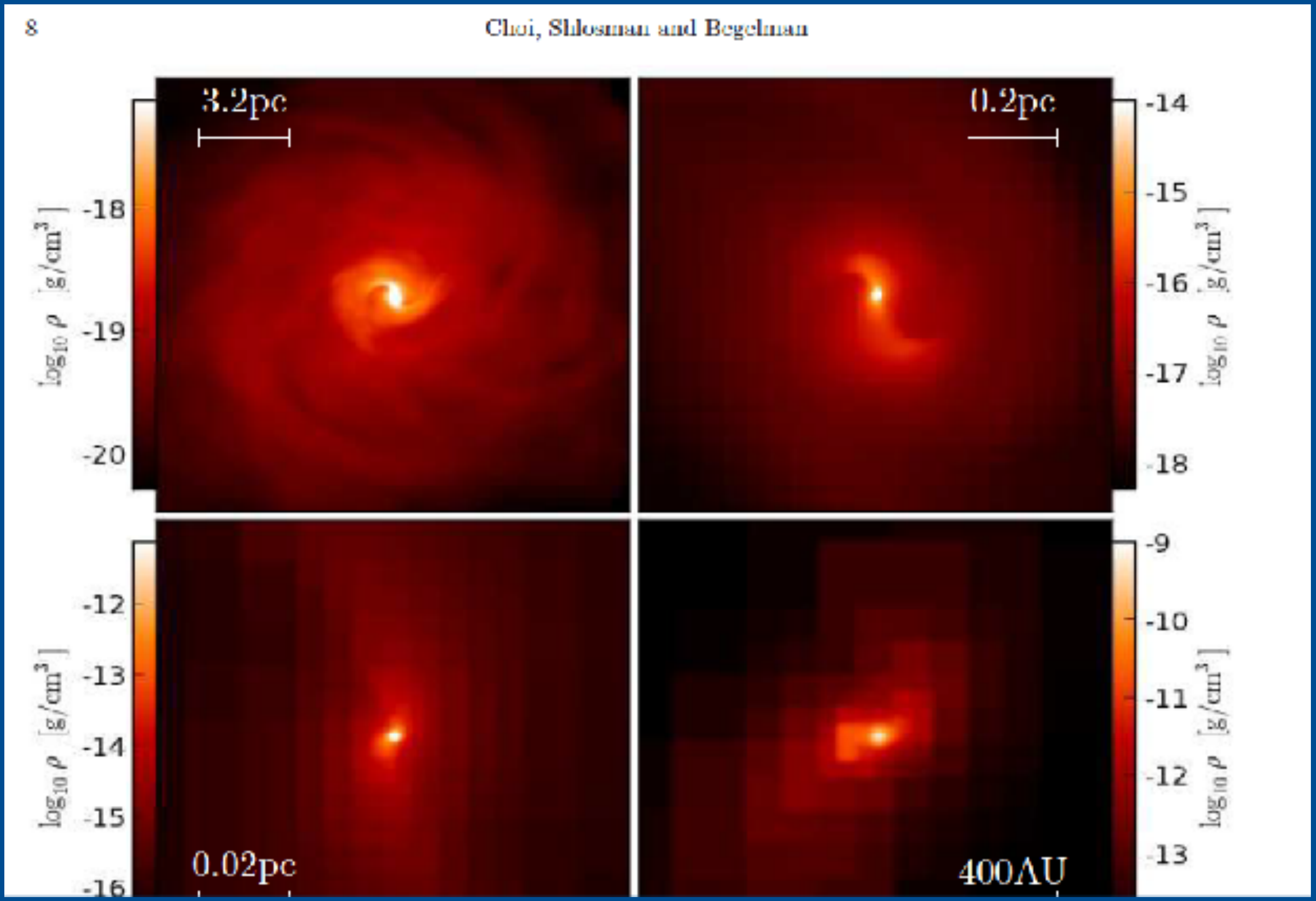
Lodato & PN 06; 07; PN 11; Regan+ 12; 14; Latif & Ferrara 16

Massive BH seed formation simulations





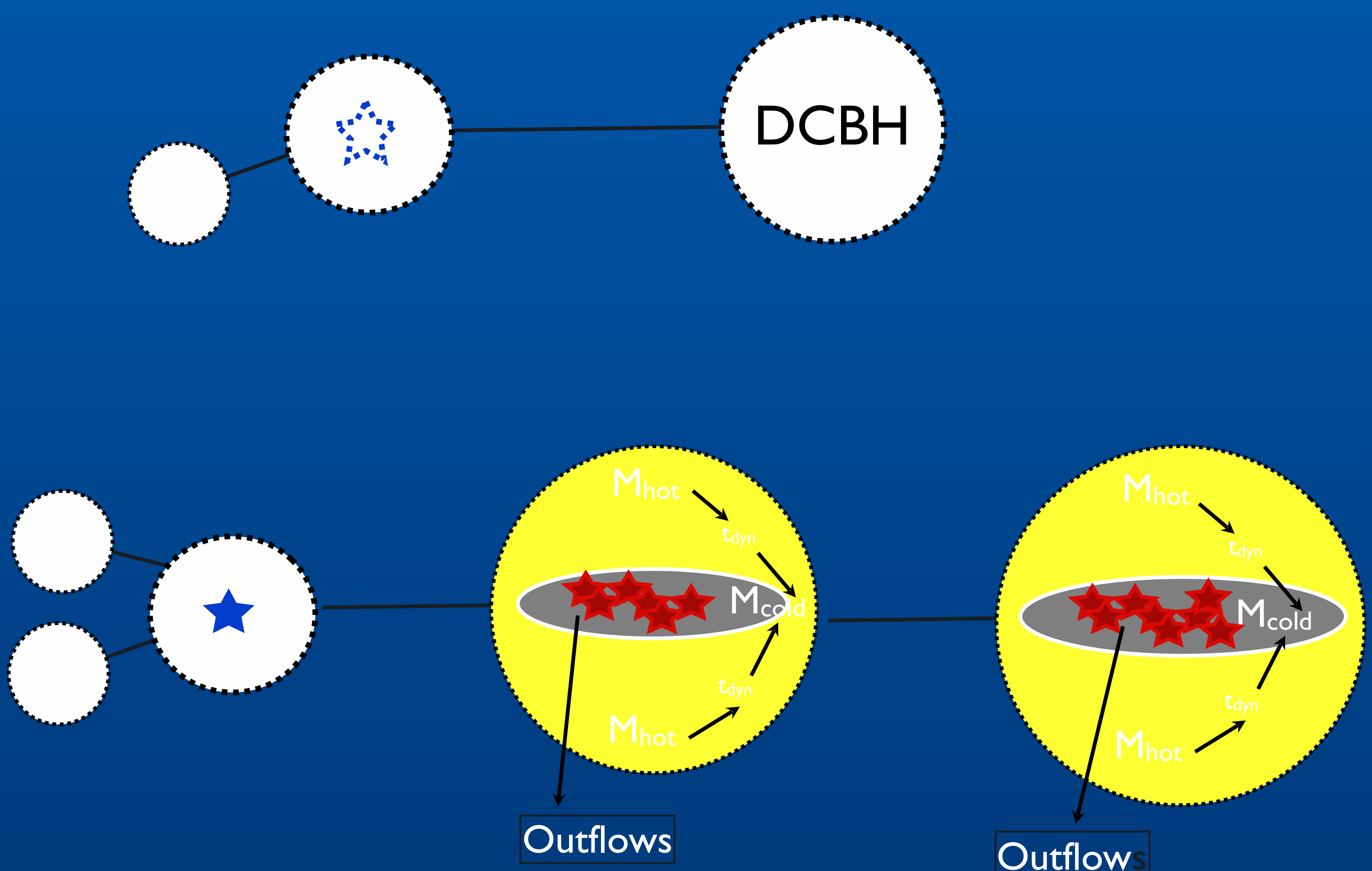
Regan+ 08; 13; Hirano+ 14; Davis & PN 12; Bournaud+; Habuzit+ 14, 15; 13; Hirano+ 14



Choi, Shlosman & Begelman 13

OPTIMAL SITES FOR DCBH FORMATION

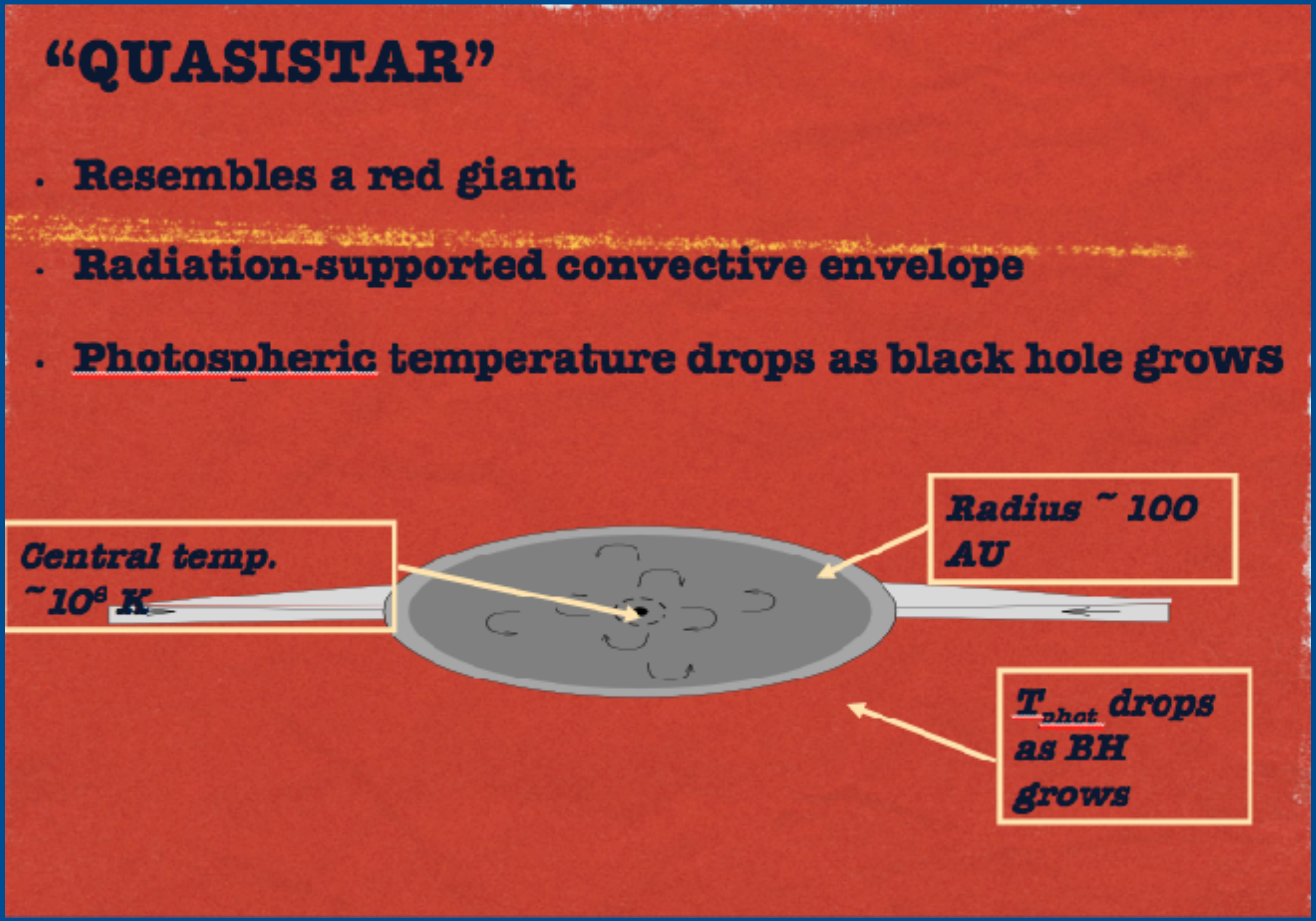
Low spin DM halos; satellite halos; Lyman-Werner radiation from nearby halos with star formation to dissociate mol H and prevent fragmentation



Agarwal+ 15

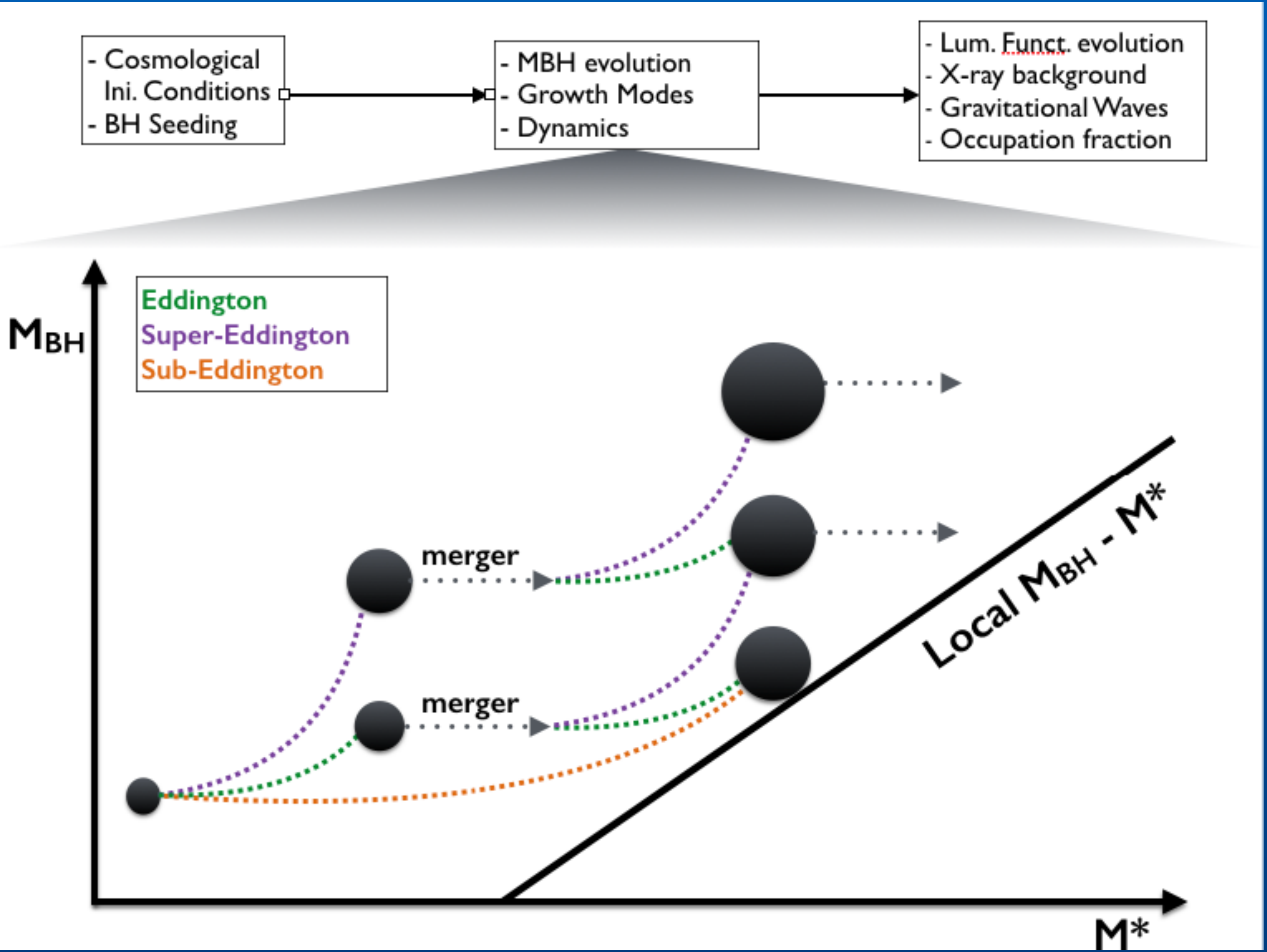
Direct Collapse Black Holes

- Metal Free gas in atomic cooling halos ($T_{vir}>10^4$ K): predominantly atomic H Avoid fragmentation into stars: high densities Prevent H_2 formation (molecular hydrogen) in central region of halo; High level of LW radiation to dissociate H_2; SMS/Quasi-star followed by black hole of mass ~ 10^4 $M_{solar;}$ N ~1% of ~ 10^7-10^8 M_{sun} DM halos at z=15 form such seeds



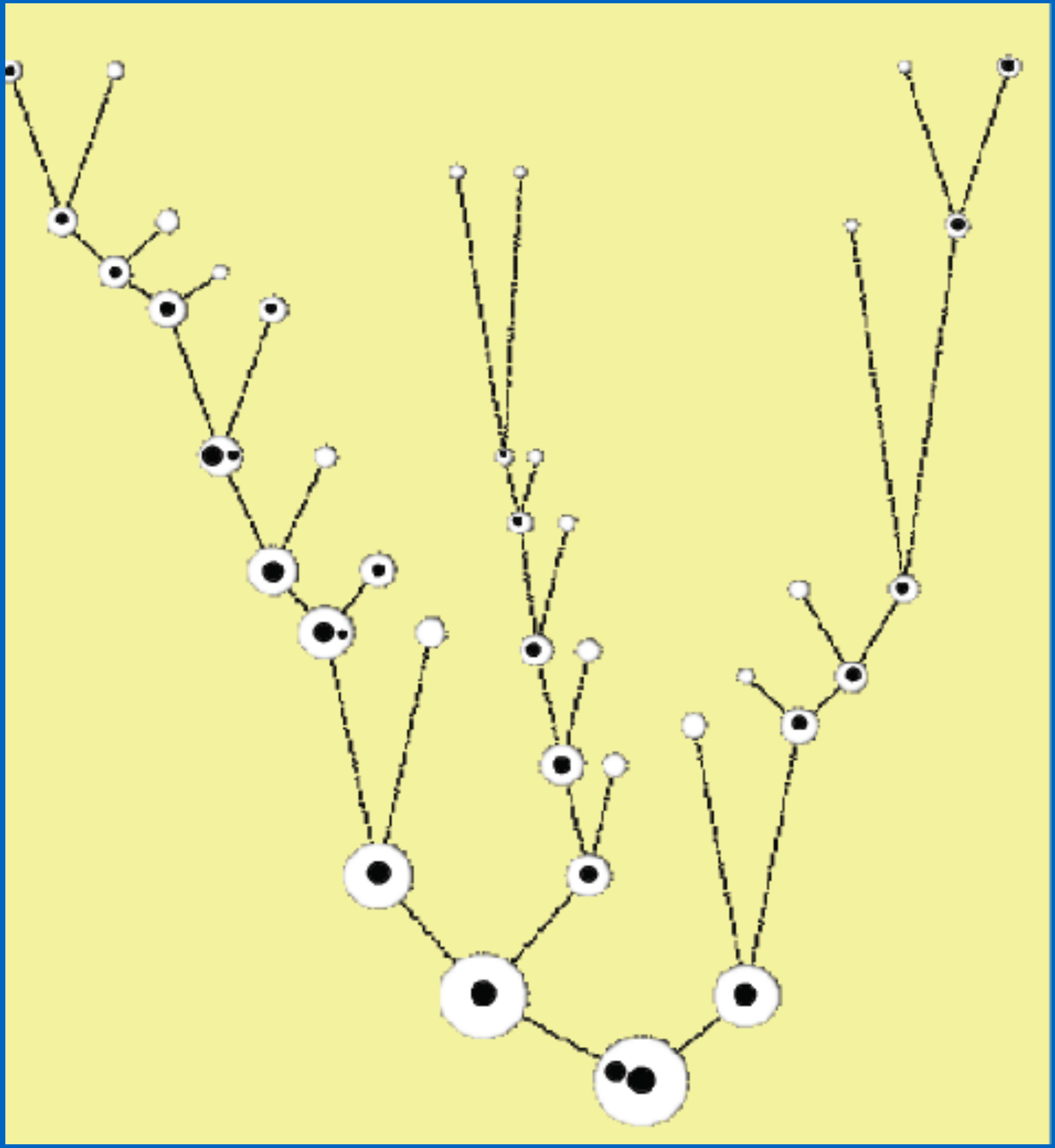
Oh & Haiman 2002; Bromm & Loeb 2003; Begelman et al. 2006; Lodato & PN 2006; 2007, Spaans & Silk 2006; Latif+; Johnson+

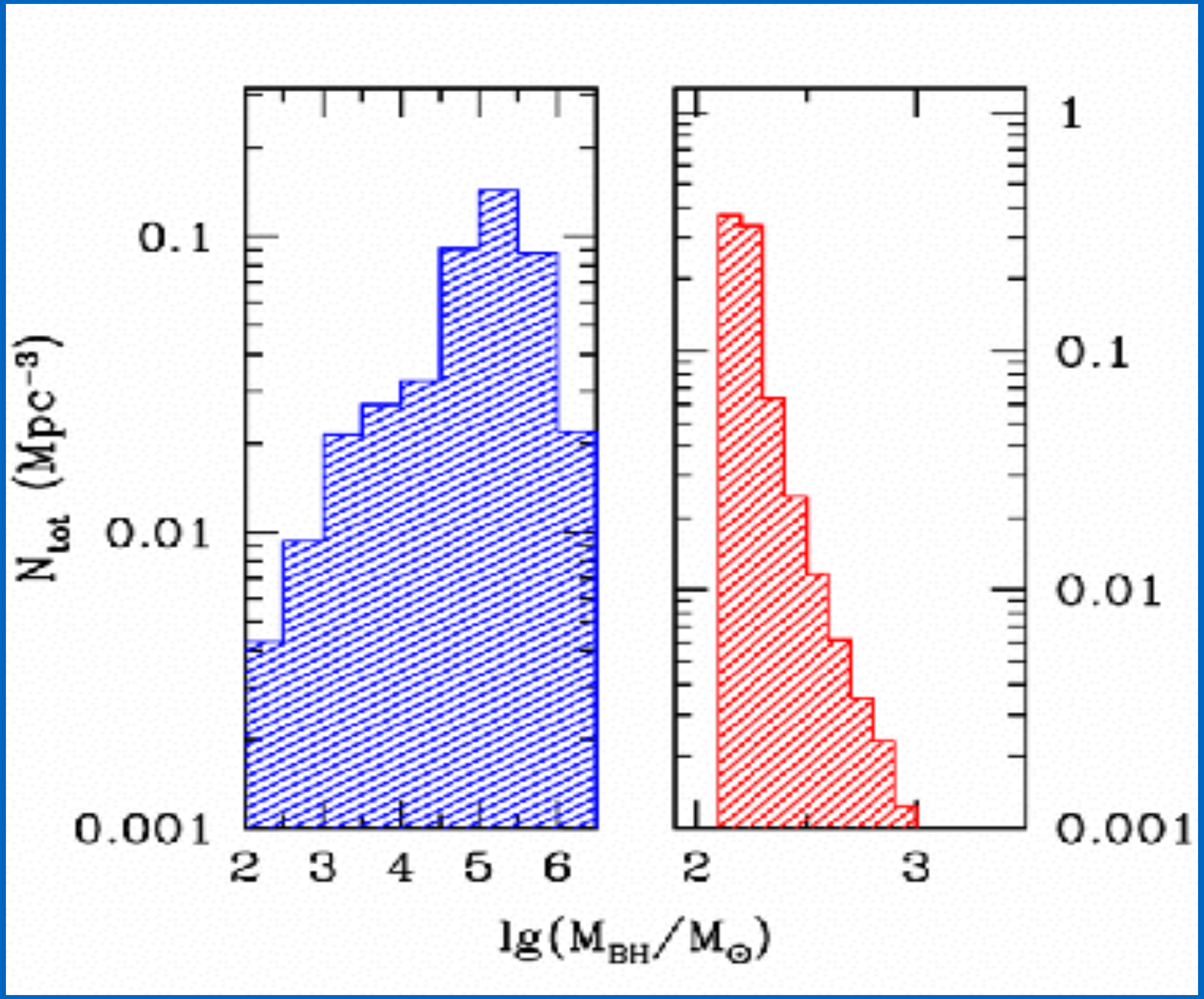
Tracing the growth history of black holes in the universe



Ricarte, PN+ 16

Merger induced accretion + CDM merger trees + BH seeds





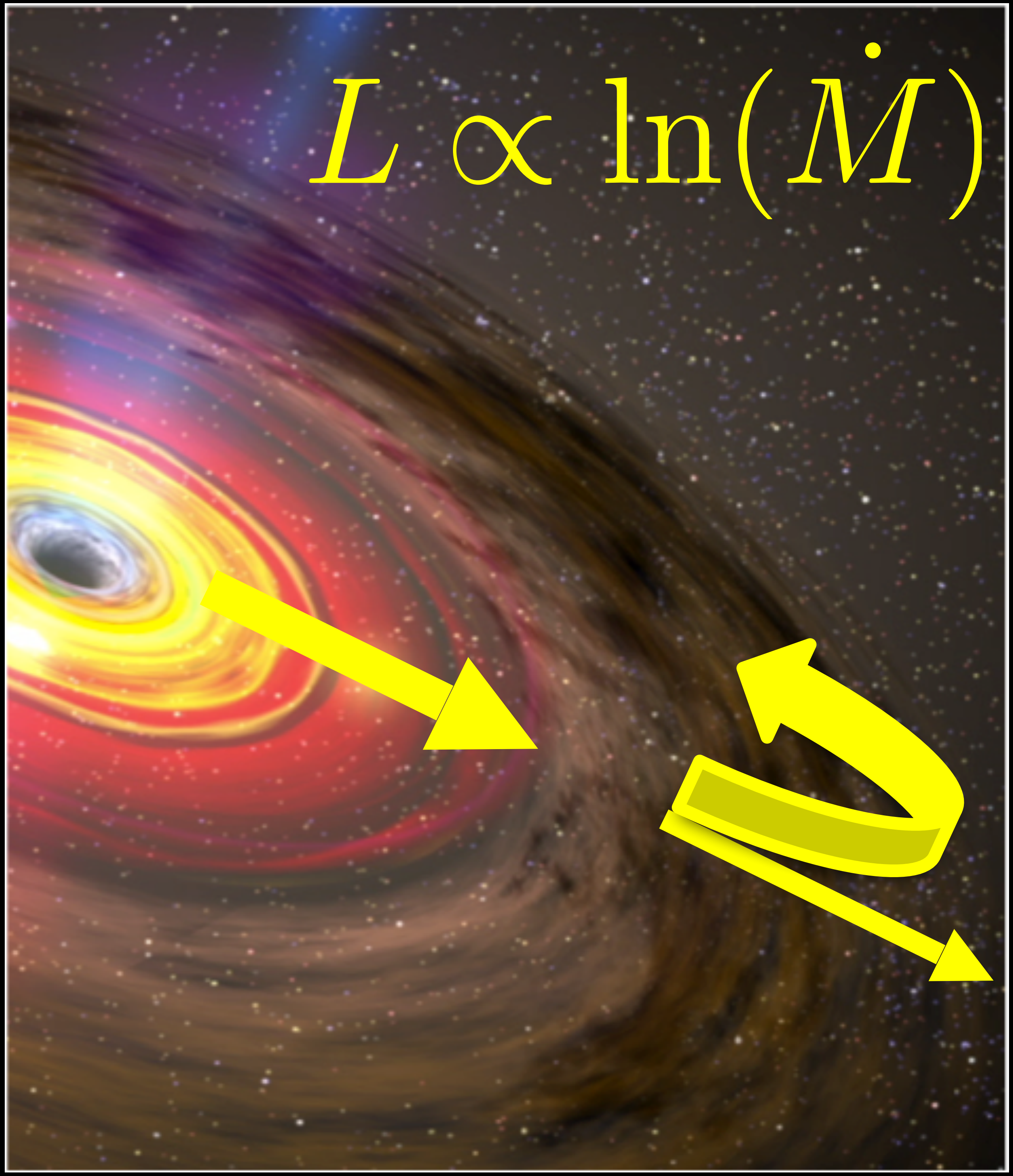
DCBH SEED PEAK MASS

Volonteri, Lodato & PN 08

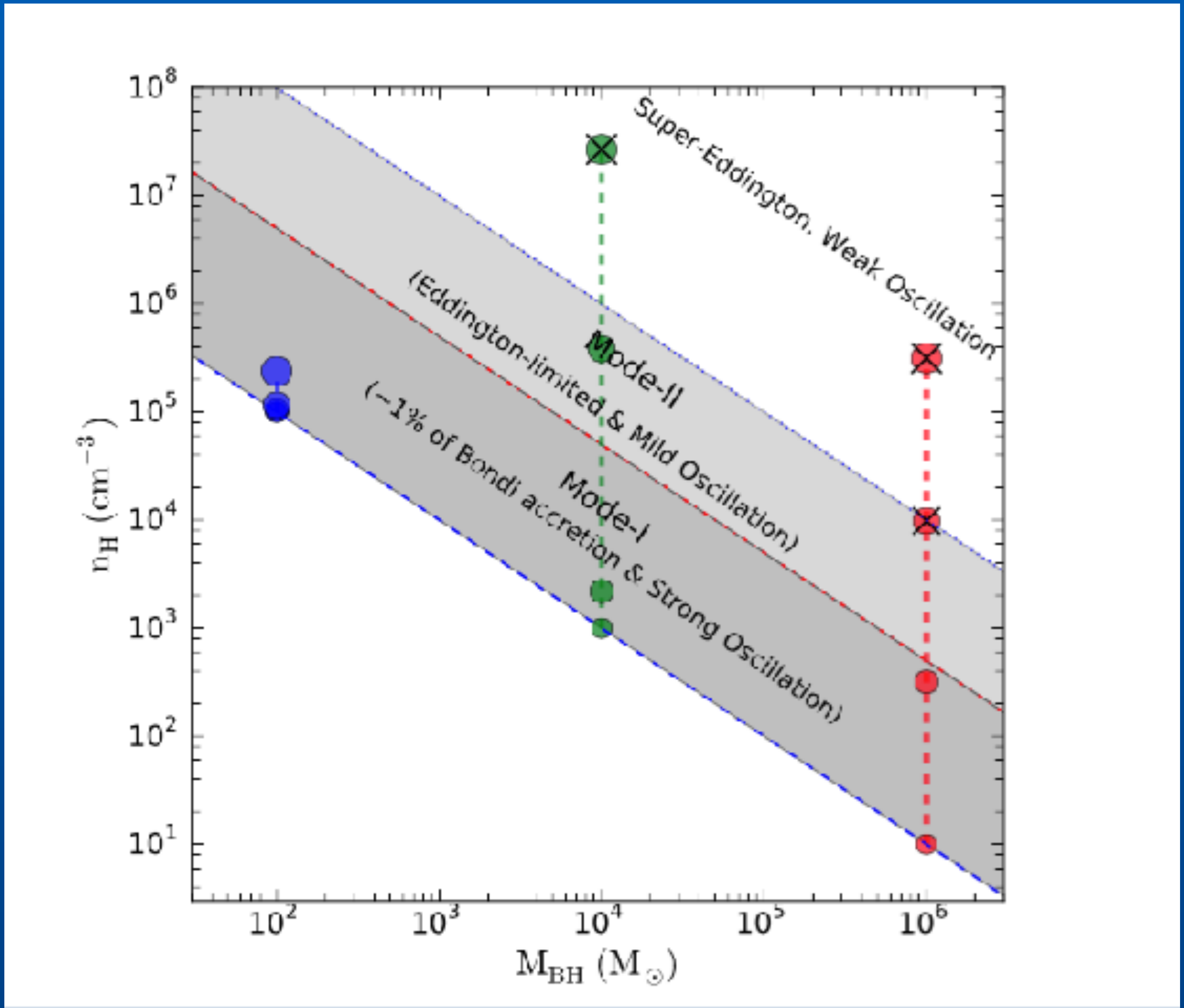
Standard Accretion

Slim Disk Accretion

Steady trickle throughout cosmic time



Understanding what limits growth rates by accretion



Growth is faster for larger black hole masses

FEEDBACK LIMITED MODE

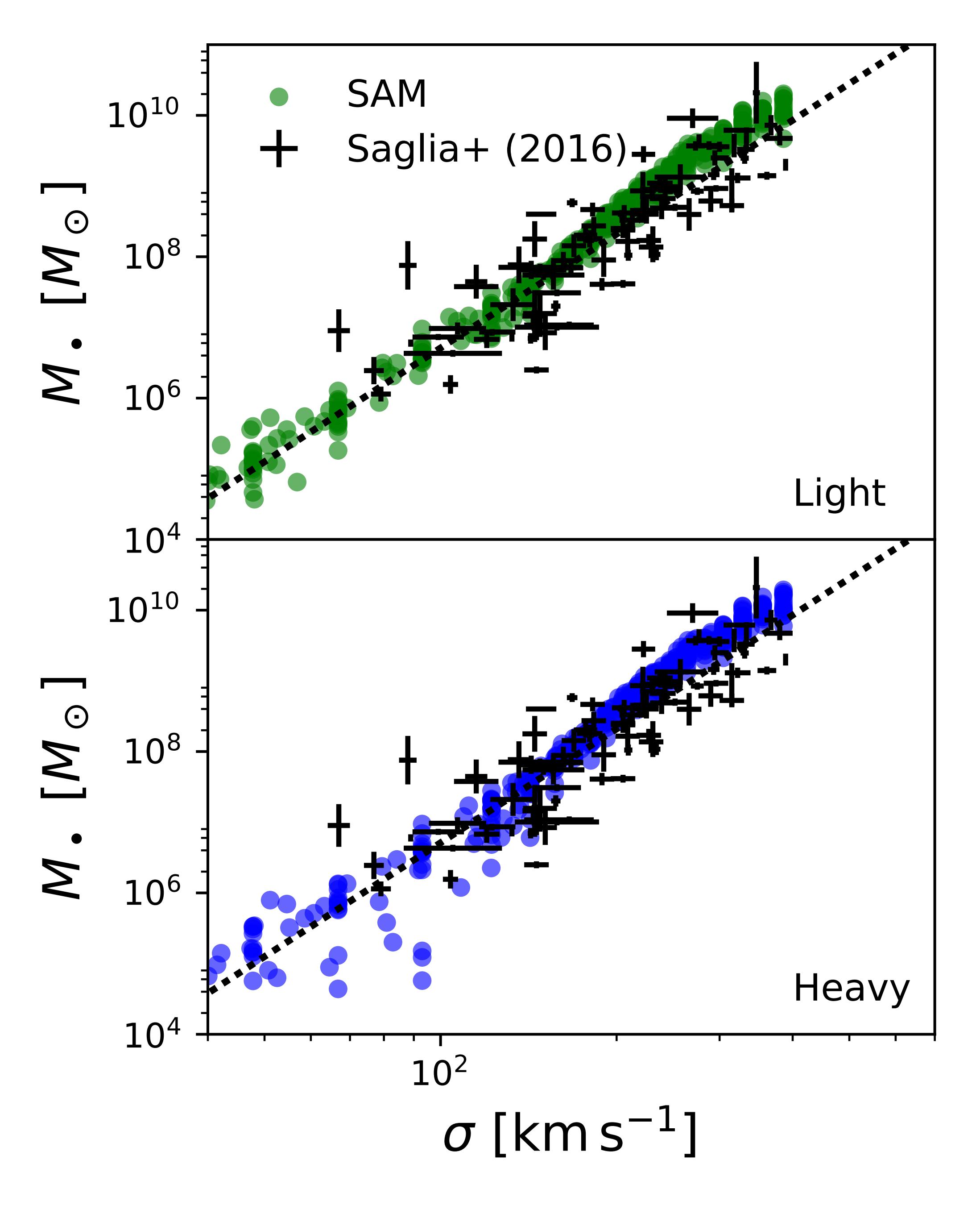
Inefficient growth, outflows,
~ 15% of gas accreted
LOW MASS SEEDS

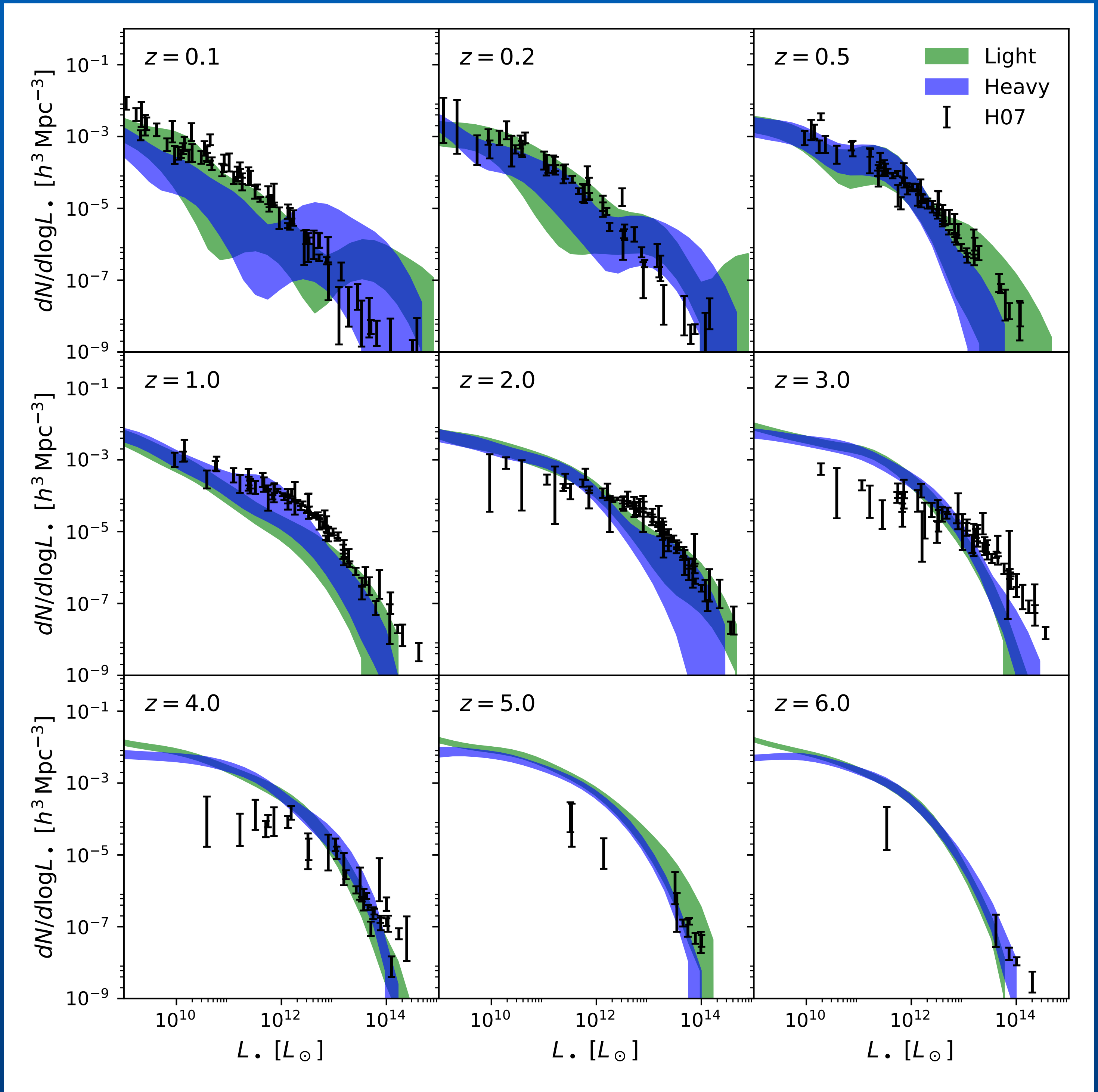
GAS SUPPLY LIMITED MODE

Super-Eddington growth, outflows unimportant, low radiative efficiency
~ 80% of gas accreted
MASSIVE SEEDS

Alexander & PN 14; Park+15; Pacucci+15

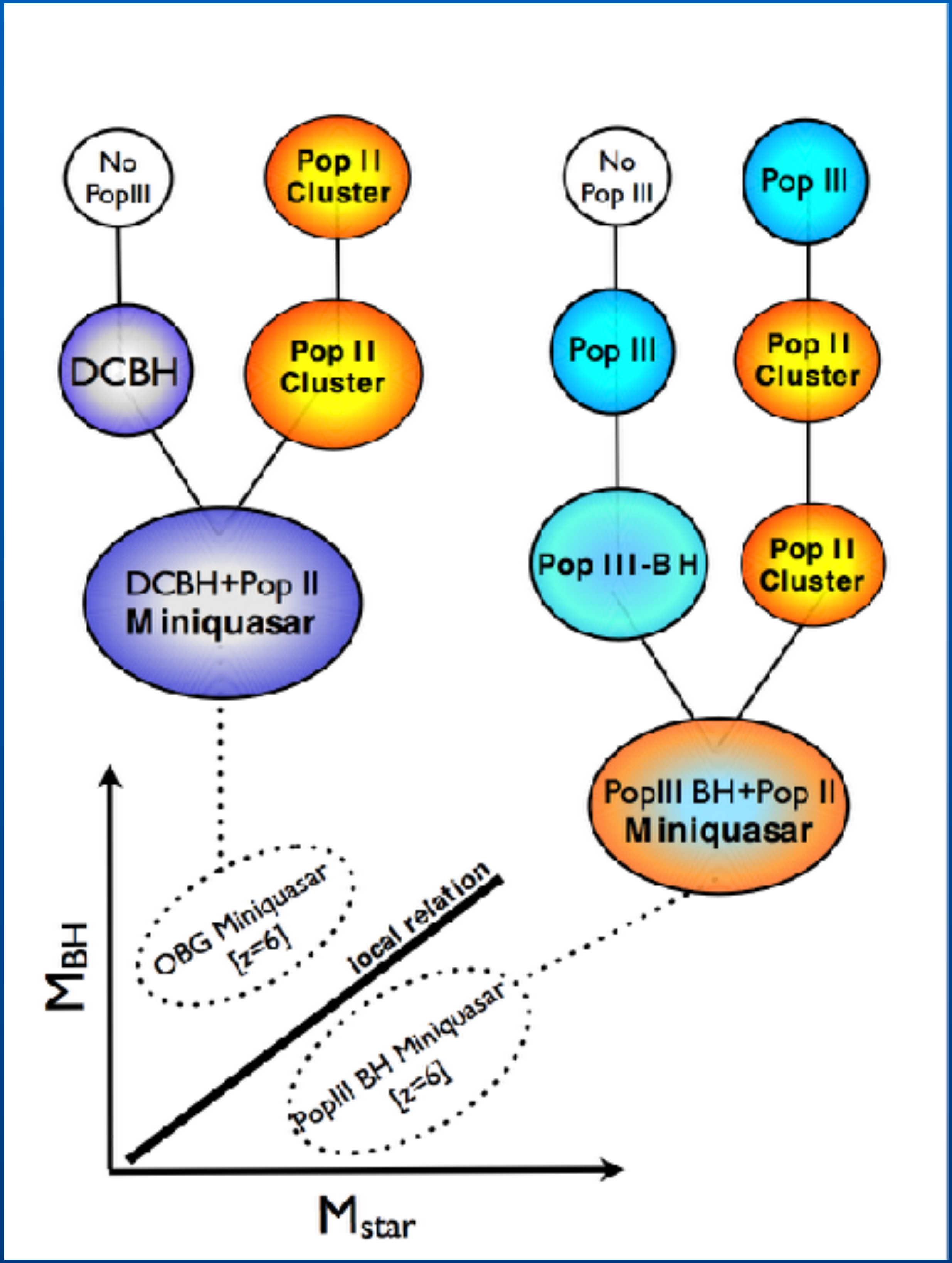
Model including stochastic accretion, flickering, mass cap





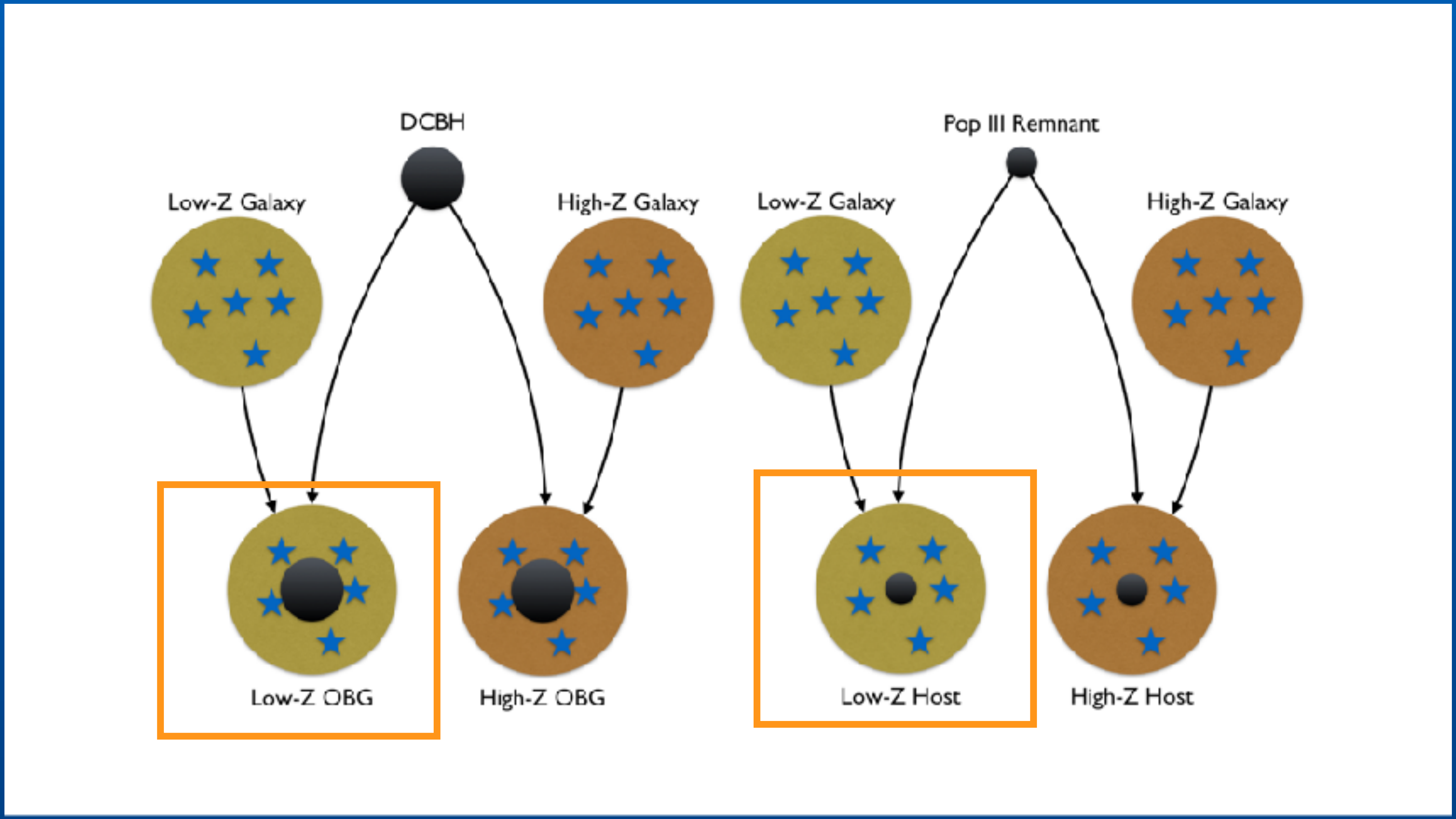
Ricarte & PN 17; Natarajan+ 17; Volonteri & PN 09

DIRECT COLLAPSE BHs AND THEIR OBESE BH HOST GALAXIES



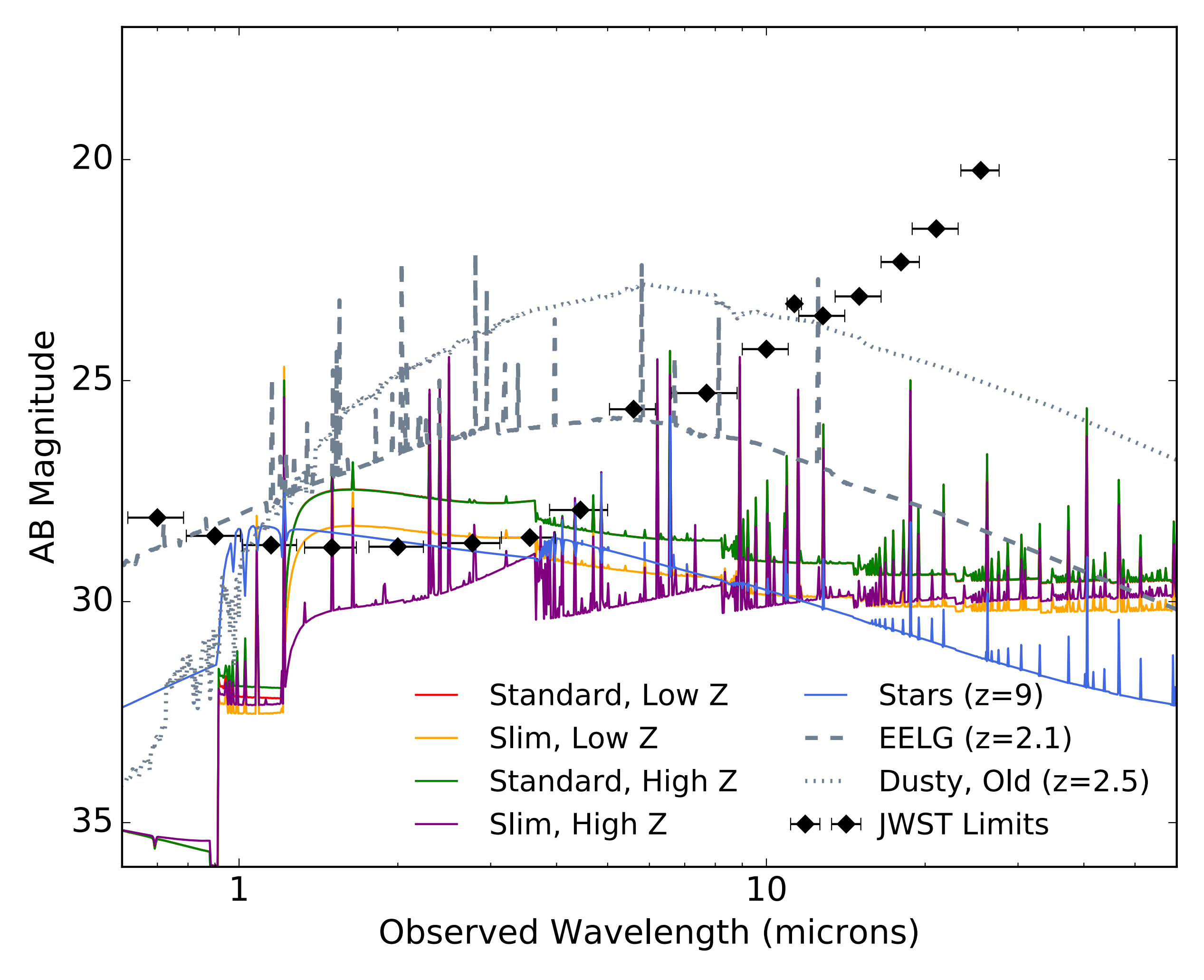
Agarwal+ 12; 14

SCHEMATIC OF HIGH-Z JWST SOURCES



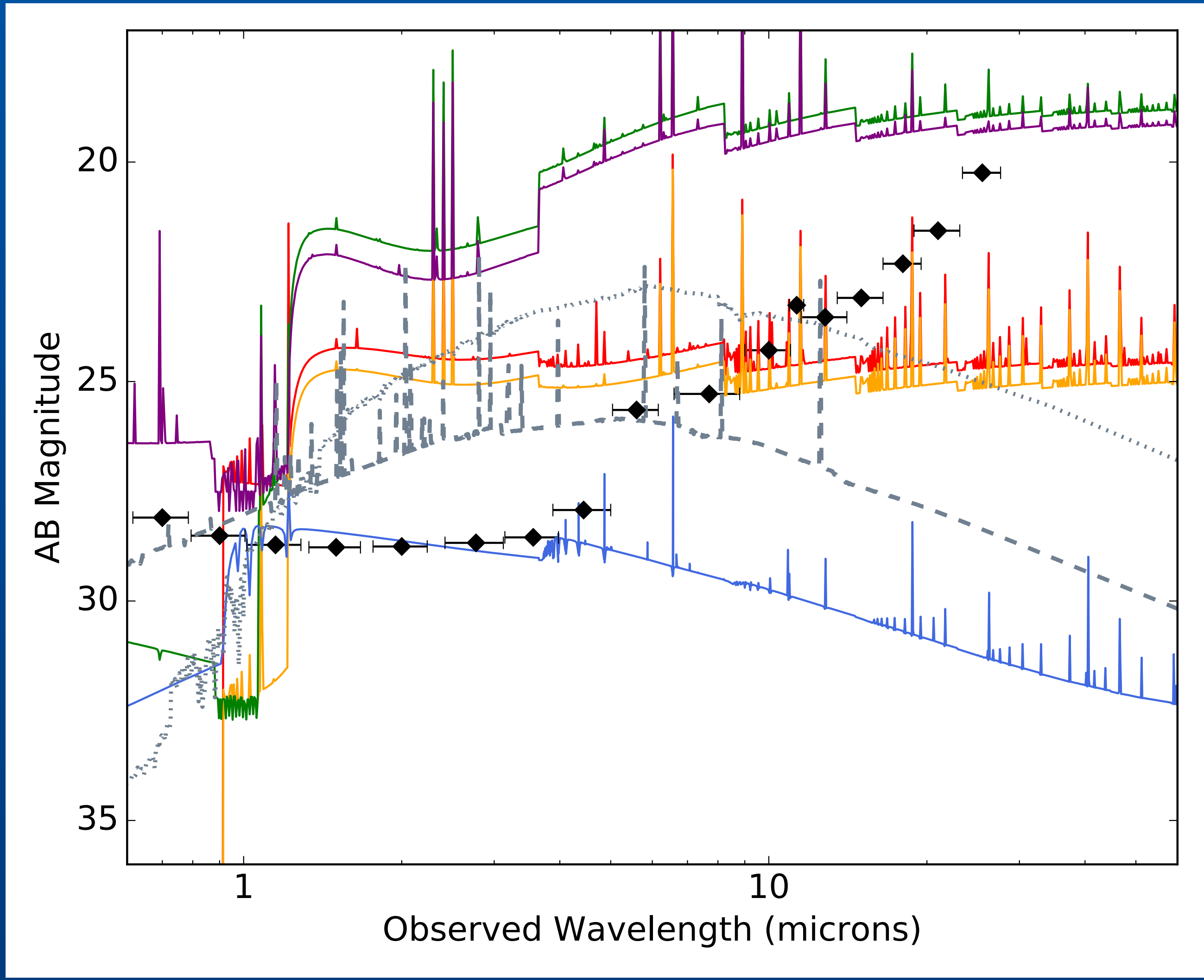
PN, Pacucci, Ferrara, Agarwal+ 16

DISCRIMINATING LIGHT & MASSIVE SEEDS WITH JWST

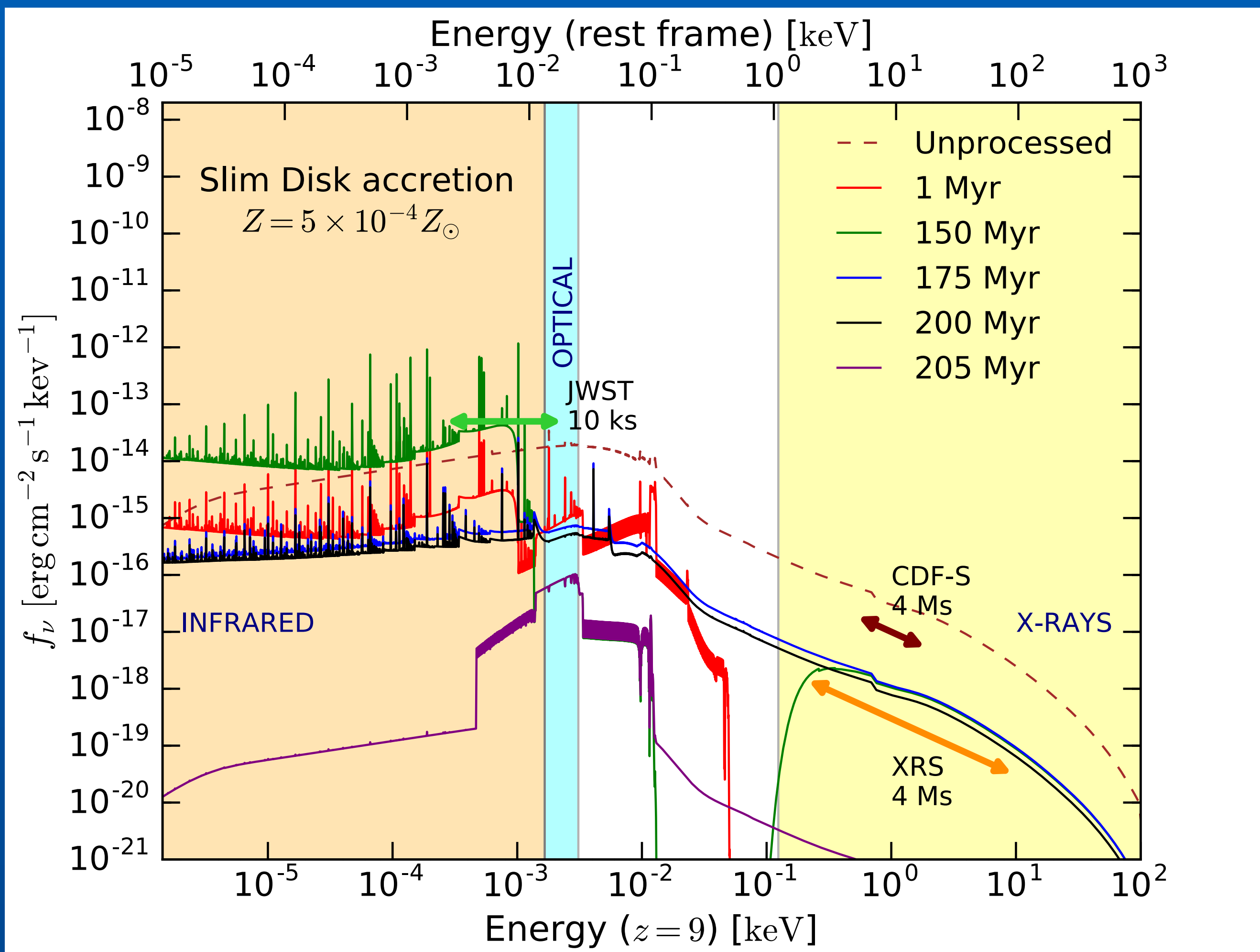


LIGHT SEEDS

MASSIVE SEEDS

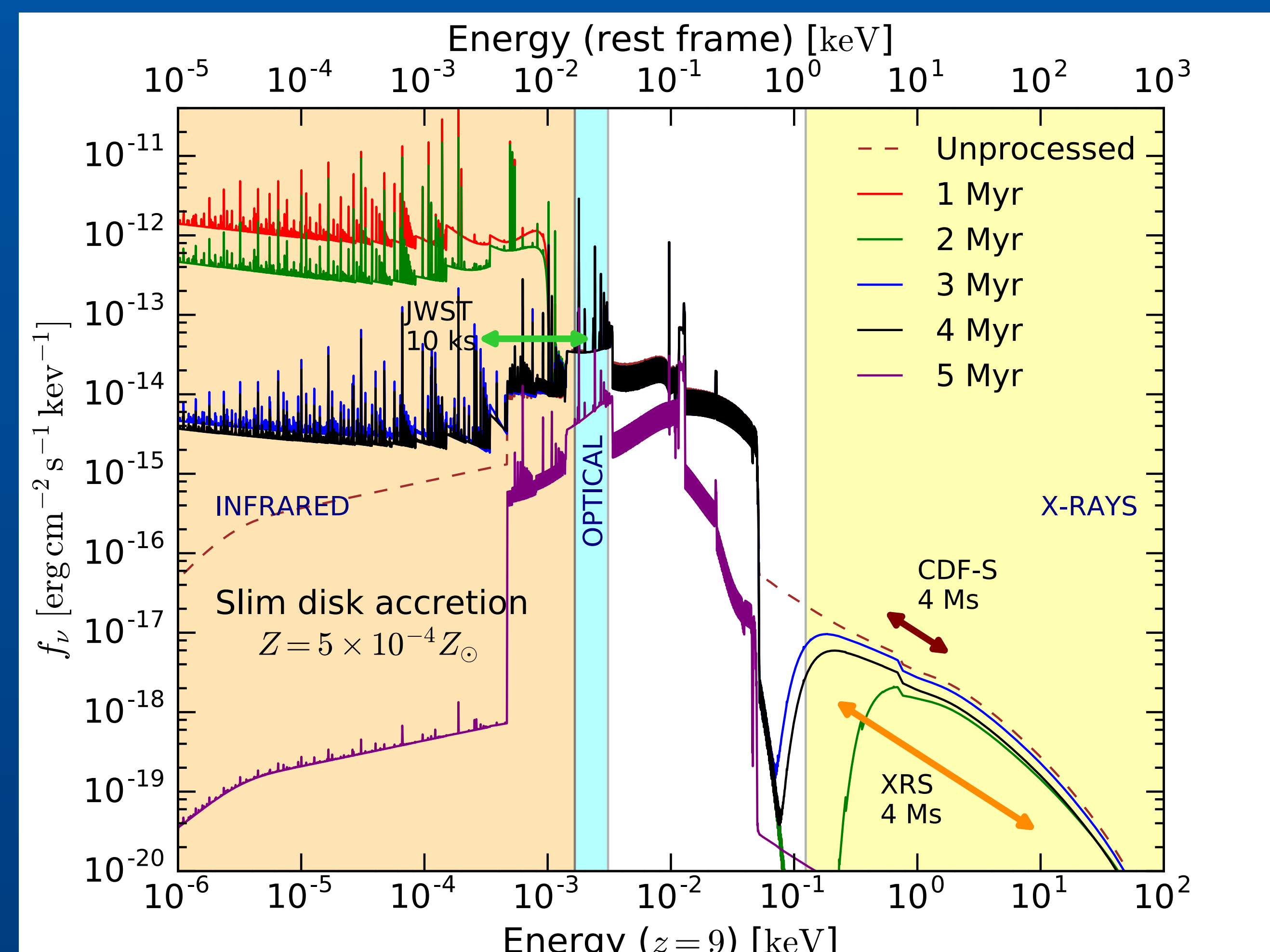


MULTI-WAVELENGTH SPECTRAL PREDICTIONS



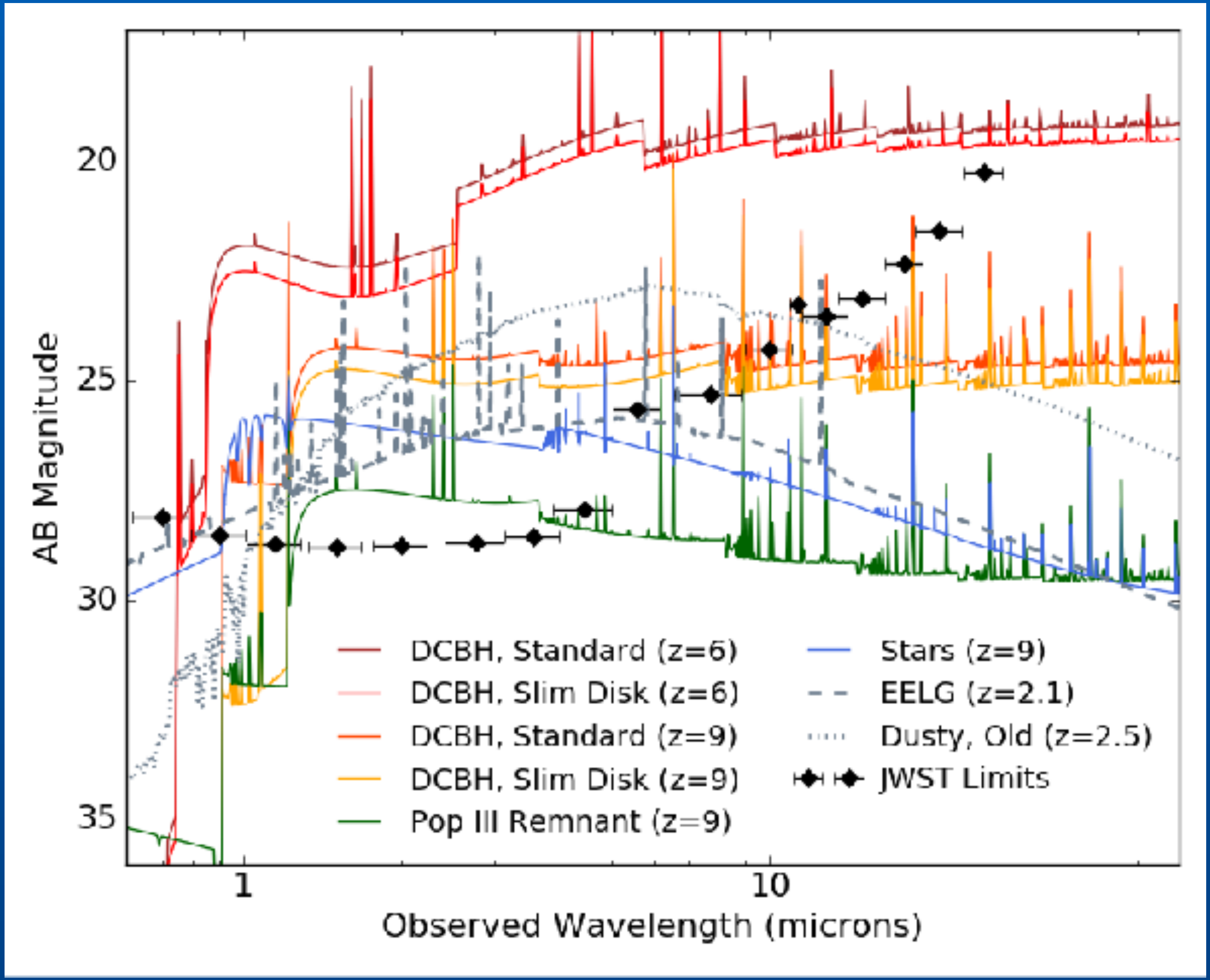
Pop III SEED + STELLAR COMPONENT SLIM DISK

DCBH SEED + STELLAR COMPONENT (OBG) SLIM DISK



Pacucci+ 15; 16; PN 16; PN, Pacucci, Ferrara, Agarwal+ 16

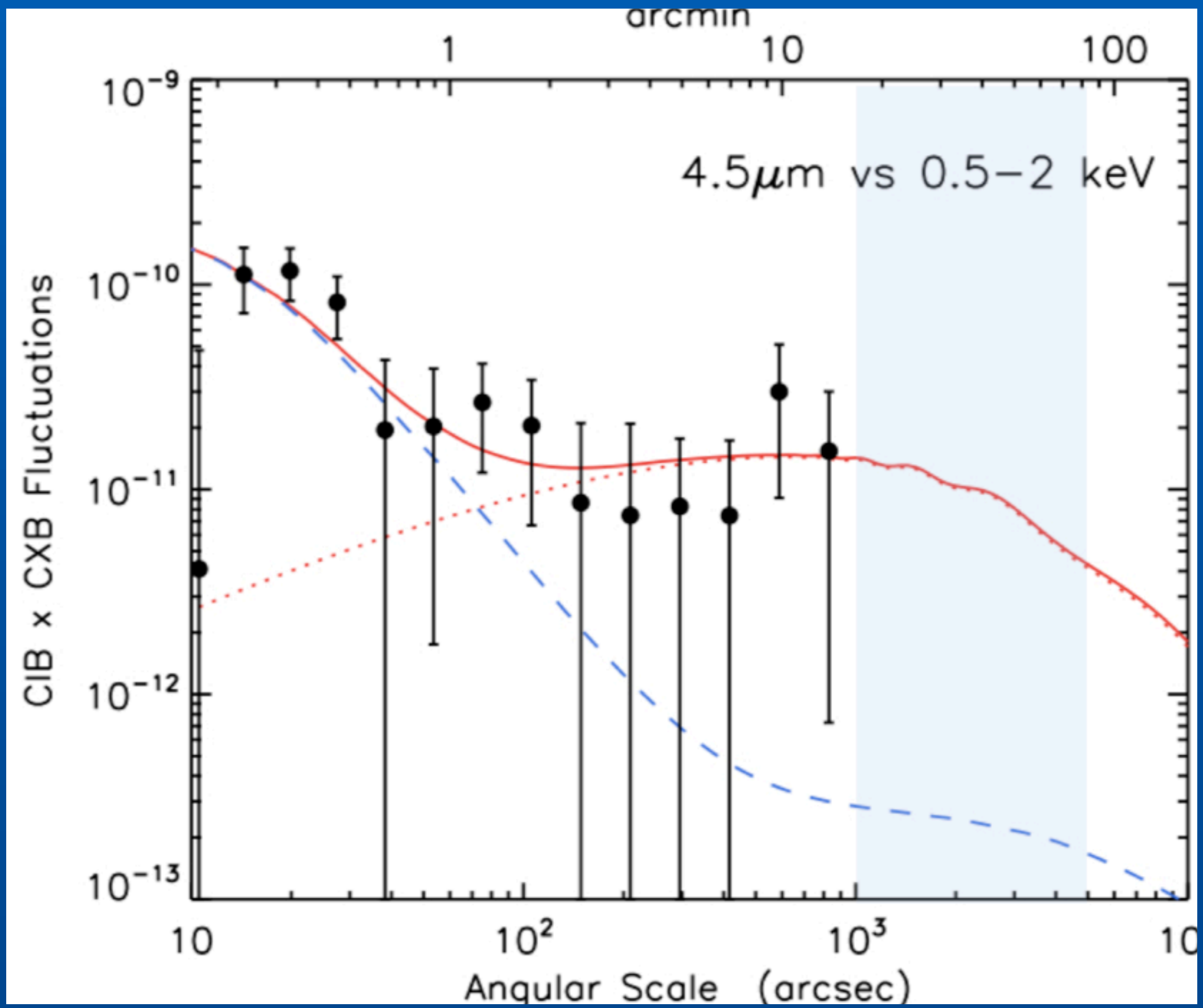
OBSERVATIONAL SIGNATURES OF DCBHs

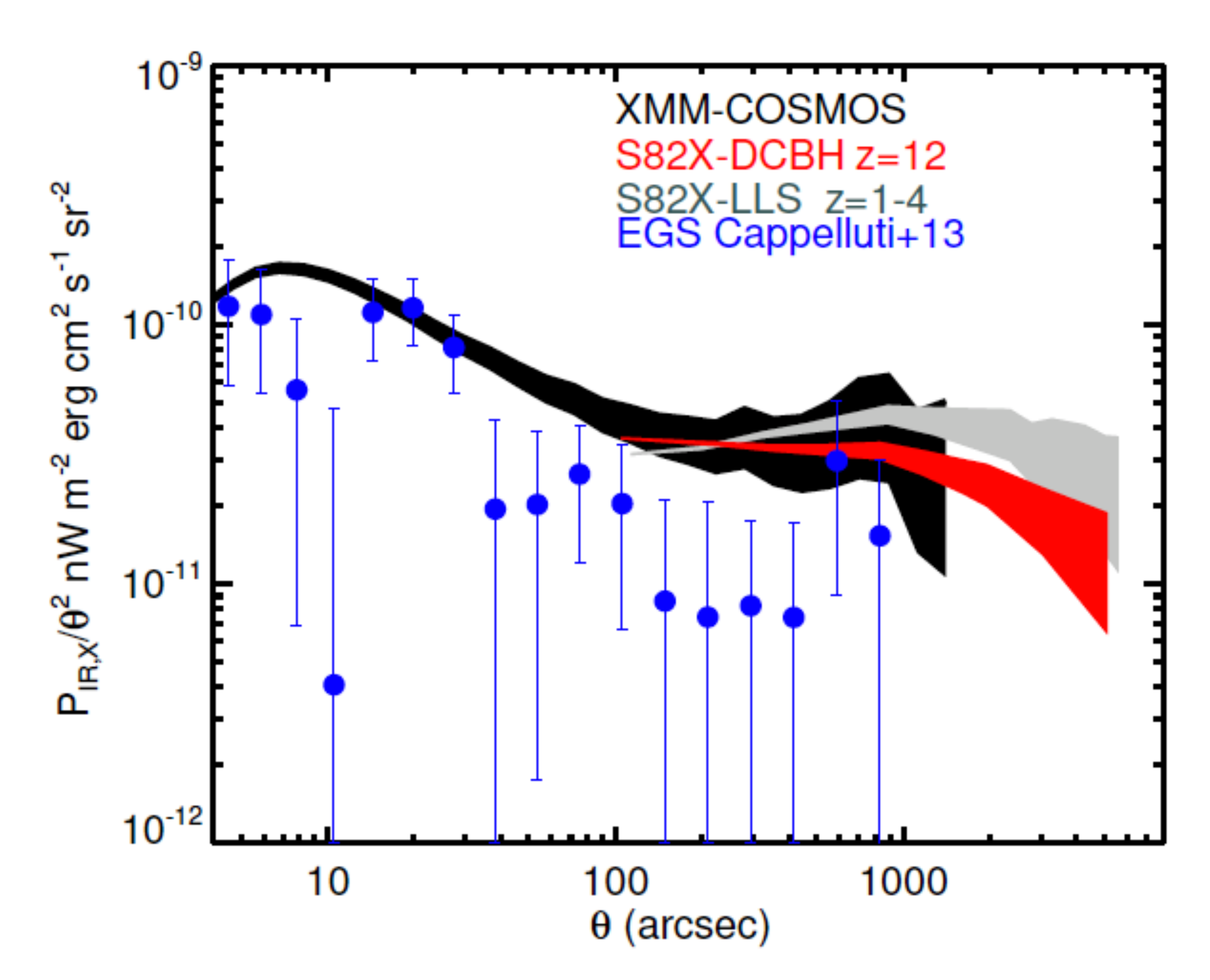


~ JWST spectra for growing DCBHs
Discriminant slope and amplitude between 1 - 10 microns

PN, Pacucci, Ferrara, Agarwal+ 16

WIDE-FIELD SURVEYS IN X-RAY & IR will reveal large scale high-z component





LARGE-SCALE CIRB EXCESS

Cappelluti+ 13; 17

STATUS REPORT OF WHERE WE ARE & OPEN QUESTIONS

Masses of initial BH seeds
Early accretion history of seed BHs
Contribution to Re-ionization
Observational signatures of Super-Eddington flows
Importance of mergers
Detection of signature of mergers – gravitational waves
When do the correlations between BHs and their hosts get set-up